

Directed searches for continuous gravitational waves from binary systems: parameter-space metrics and optimal Scorpius X-1 sensitivity

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We derive simple analytic expressions for the (coherent and semi-coherent) phase metrics of continuous-wave sources in low-eccentricity binary systems, for the two regimes of long and short segments compared to the orbital period. The resulting expressions correct and extend previous results found in the literature. We present results of extensive Monte-Carlo studies comparing metric mismatch predictions against the measured loss of detection statistic for binary parameter offsets. The agreement is generally found to be within $\sim 10\% - 30\%$. As an application of the metric template expressions, we estimate the optimal achievable sensitivity of an Einstein@Home directed search for Scorpius X-1, under the assumption of sufficiently small spin wandering. We find that such a search, using data from the upcoming advanced detectors, would be able to beat the torque-balance level [1, 2] up to a frequency of $\sim 500 - 600$ Hz, if orbital eccentricity is well-constrained, and up to a frequency of $\sim 160 - 200$ Hz for more conservative assumptions about the uncertainty on orbital eccentricity.

I. INTRODUCTION

Continuous gravitational waves (CWs) are a promising class of signals for the second-generation detectors currently under construction: advanced LIGO (aLIGO) [3], advanced Virgo [4] and KAGRA [5]. These signals would be emitted by spinning neutron stars (NSs) subject to non-axisymmetric deformations, such as quadrupolar deformations (“mountains”), unstable oscillation modes (e.g. r-modes) or free precession. For a general review of CW sources and search methods, see for example [6].

A particularly interesting type of potential CW sources are NSs in low-mass X-ray binaries (LMXBs), with Scorpius X-1 being its most prominent representative [7]. The accretion in these systems would be expected to have spun up the NSs to the maximal rotation rate of $\gtrsim 1000$ Hz [8]. All the observations performed to date, however, show they only spin at several hundred Hz, so there seems to be something limiting the accretion-induced spin-up. A limiting mechanism that has been suggested is the emission of gravitational waves [1, 2, 9], which would result in steady-state CW emission where the accretion torque is balanced by the radiated angular momentum. For a discussion of alternative explanations, see [2, 9]. The resulting torque-balance CW amplitude increases with the observed X-ray flux (independent of the distance of the system, e.g. see Eq.(4) in [2]), therefore the X-ray brightest LMXB Scorpius X-1 is often considered the most promising CW source within this category [7, 10].

Several searches for CW signals from Scorpius X-1 have been performed (without any detections) on data from

initial LIGO [11–13], and several new pipelines have been developed and are currently being tested in a Scorpius X-1 Mock Data Challenge (MDC) [10]. So far, all current methods fall into either one of two extreme cases: highly coherent with a short total time baseline (6 hours in [11], 10 days in [13]), or highly incoherent with a long total baseline but very short (\sim hours) coherent segments [12, 14, 15].

In order to increase sensitivity beyond these methods, and to be able to effectively absorb large amounts of computing power (such as those provided by Einstein@Home [16, 17]), it is necessary to extend the search approach into the realm of general long-segment semi-coherent methods, by stacking coherent \mathcal{F} -statistic segments [18–20]. Such methods have already been employed over the past years for both directed and all-sky searches for CWs from isolated NSs [17, 21, 22], but they have not yet been extended to the search for CWs from binary systems.

Key ingredients required for building a semi-coherent “StackSlide” search are the coherent and semi-coherent parameter-space metrics [18, 23–25]. The coherent binary CW metric was first analyzed in [26], and this was further developed and extended to the semi-coherent case in [27]. Here we will largely follow the approach of [27], considering only low-eccentricity orbits, and focusing on two different regimes of either long coherent segments, or short segments compared to the orbital period, both of which admit simple analytic results.

Plan of this paper

Sec. II provides a general introduction of the concepts and notation of semi-coherent StackSlide methods, parameter-space metric and template banks. Secs. III–V build on the work of [27], rectifying some of the results and extending them to the case of a general orbital refer-

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ence time. In Sec. VI we present the results of extensive Monte-Carlo software-injection tests comparing the metric predictions against the measured loss in statistics due to parameter-space offsets. Sec. VII applies the theoretical template-bank counts to compute optimal StackSlide setups for a directed search for Scorpius X-1 and estimates the resulting achievable sensitivity. Finally, Sec. VIII presents concluding remarks.

Notation. Throughout the paper we denote a quantity Q as \tilde{Q} when referring to the coherent case, and as \hat{Q} when referring to the semi-coherent case.

II. BACKGROUND

The strain h in a given detector due to a continuous gravitational wave is a scalar function $h(t; \mathcal{A}, \lambda)$, where t is the time at the detector. The set of signal parameters \mathcal{A} denotes the four *amplitude parameters*, namely the overall amplitude h_0 , polarization angles $\cos \iota$ and ψ , and the initial phase ϕ_0 . The set of *phase-evolution parameters* λ consist of all the remaining signal parameters affecting the time-evolution of the CW phase at the detector, notably the signal frequency f , its frequency derivatives (also known as “spindown” terms), sky-position of the source, and binary orbital parameters. We will look in more detail at the phase model in Sec. III.

The observed strain $x(t)$ in the detector affected by additive Gaussian noise $n(t)$ can be written as $x(t) = n(t) + h(t; \mathcal{A}, \lambda)$, and the *detection problem* consists in distinguishing the (pure) “noise hypothesis” of $h_0 = 0$ from the “signal hypothesis” $h_0 > 0$.

A. Coherent detection statistic

As shown in [28], matched-filtering the data $x(t)$ against a template $h(t; \mathcal{A}, \lambda)$, and analytically maximizing over the unknown amplitude parameters \mathcal{A} , results in the coherent statistic $\tilde{\mathcal{F}}(x; \lambda)$, which only depends on the data x and on the template phase-evolution parameters λ . This statistic follows a (non-central) χ^2 distribution with 4 degrees of freedom, and a non-centrality parameter $\tilde{\rho}^2(\mathcal{A}, \lambda_s; \lambda)$, which depends on the signal parameters $\{\mathcal{A}, \lambda_s\}$ and the template parameters λ . The quantity $\tilde{\rho}$ (which depends linearly on h_0) is generally referred to as the signal-to-noise ratio (SNR) of the coherent $\tilde{\mathcal{F}}$ -statistic. The expectation value of this statistic over noise realizations is given by [28]

$$E \left[2\tilde{\mathcal{F}}(x; \lambda) \right] = 4 + \tilde{\rho}^2(\mathcal{A}, \lambda_s; \lambda). \quad (1)$$

In the case of unknown signal parameters $\lambda_s \in \mathcal{P}$ within some parameter space \mathcal{P} , the number of required template λ_t that need to be searched is typically impractically large. The maximal achievable sensitivity by this coherent statistic $\tilde{\mathcal{F}}(x; \lambda)$ is therefore severely limited [25]

by the required computing cost. It was shown that *semi-coherent* statistics generally result in better sensitivity at equal computing cost (e.g. see [18, 20]).

B. Semi-coherent detection statistic

Here we focus on one particular semi-coherent approach, sometimes referred to as StackSlide, which consists of dividing the total amount of data T_{obs} into N segments of duration ΔT , such that $T_{\text{obs}} = N \Delta T$ (in the ideal case of gapless data). The coherent statistic $\tilde{\mathcal{F}}_\ell(x; \lambda)$ is then computed over all segments $\ell = 1 \dots N$, and combined incoherently by summing¹

$$\hat{\mathcal{F}}(x; \lambda) \equiv \sum_{\ell=1}^N \tilde{\mathcal{F}}_\ell(x; \lambda). \quad (2)$$

This statistic follows a (non-central) χ^2 -distribution with $4N$ degrees of freedom and a non-centrality parameter $\hat{\rho}^2(\mathcal{A}, \lambda_s; \lambda)$, which is found to be given by

$$\hat{\rho}^2 = \sum_{\ell=1}^N \tilde{\rho}_\ell^2, \quad (3)$$

in terms of the per-segment coherent SNRs $\tilde{\rho}_\ell$. Note that (contrary to the coherent case) $\hat{\rho}$ can *not* be regarded as a semi-coherent SNR. The expectation for $\hat{\mathcal{F}}$ is

$$E \left[2\hat{\mathcal{F}}(x; \lambda) \right] = 4N + \hat{\rho}^2(\mathcal{A}, \lambda_s; \lambda). \quad (4)$$

C. Template banks and metric mismatch

In order to systematically search a parameter space \mathcal{P} using a statistic $\mathcal{F}(x; \lambda)$ (which here can refer either to the coherent $\tilde{\mathcal{F}}$ or the semi-coherent $\hat{\mathcal{F}}$), we need to select a finite sampling $\{\lambda_t\} \subset \mathcal{P}$ of the parameter space, commonly referred to as a template bank, and compute the statistic over this set of templates, i.e. $\{\mathcal{F}(x; \lambda_t)\}$.

A signal with parameters $\lambda_s \in \mathcal{P}$ will generally not fall on an exact template, and we therefore need to characterize the loss of detection statistic \mathcal{F} as a function of the offset $\delta\lambda = \lambda - \lambda_s$ from a signal. This is generally quantified using the *expected* statistic $E[2\mathcal{F}]$ of Eqs. (1) and (4), namely (by removing the bias $4N$) as the relative loss in the non-centrality ρ^2 :

$$\mu_0 \equiv \frac{\rho^2(\mathcal{A}, \lambda_s; \lambda_s) - \rho^2(\mathcal{A}, \lambda_s; \lambda)}{\rho^2(\mathcal{A}, \lambda_s; \lambda_s)}, \quad (5)$$

¹ This form of the “ideal” StackSlide statistic is not directly used for actual searches. For reasons of computational cost, the coherent statistics across segments are computed on coarser template banks, and are combined on a fine template bank by summing across segments using interpolation on the coarse grids. This is discussed in detail in [20], but is not relevant for the present investigation.

which defines the measured \mathcal{F} -statistic *mismatch* function. This mismatch has a global minimum of $\mu_0 = 0$ at $\lambda = \lambda_s$ and is bounded within $\mu_0 \in [0, 1]$.

We now use two standard approximations: (i) Taylor-expand this up to second order in small offsets $\delta\lambda$, and (ii) neglect the dependence on the amplitude parameters \mathcal{A} (the effect of which was analyzed in detail in [29]), which leads to the well-known *phase metric* g_{ij} , namely

$$\mu_0(\mathcal{A}, \lambda_s; \lambda) \approx \mu(\lambda_s; \lambda), \quad \text{where} \quad (6)$$

$$\mu(\lambda_s; \lambda) \equiv g_{ij}(\lambda_s) \delta\lambda^i \delta\lambda^j, \quad (7)$$

with g_{ij} a positive-definite symmetric matrix, and implicit summation over repeated indices i, j is used. The quality of this approximation will be quantified in a Monte-Carlo study in Sec. VI. The *metric mismatch* μ has a global minimum of $\mu = 0$ at $\lambda = \lambda_s$ and (contrary to μ_0) is semi-unbounded, i.e. $\mu \in [0, \infty)$. Previous studies testing μ against μ_0 [29–31] for all-sky searches have shown that the phase-metric approximation Eq. (6) works reasonably well for observation times $T_{\text{obs}} \gtrsim \mathcal{O}(1 \text{ day})$ and for mismatch values up to $\mu \lesssim 0.3 - 0.5$. Above this mismatch regime, the metric mismatch μ increasingly *over-estimates* the actual loss μ_0 . For example, in [29] an empirical fit for this behavior was given as $\mu_0(\mu) \approx \mu - 0.38\mu^2$.

The metric formalism was first introduced in [23, 24] (in the context of compact binary coalescence signals), and was later applied to the CW problem in [25], where it was shown that the *coherent* phase metric \tilde{g}_{ij} can be expressed explicitly as

$$\tilde{g}_{ij}(\lambda) = \langle \partial_i \phi(\lambda) \partial_j \phi(\lambda) \rangle - \langle \partial_i \phi(\lambda) \rangle \langle \partial_j \phi(\lambda) \rangle, \quad (8)$$

where $\partial_i \phi \equiv \partial \phi / \partial \lambda^i$, with $\phi(t; \lambda)$ the CW signal phase, and $\langle \dots \rangle$ denotes the time average over the coherence time T , i.e. $\langle Q \rangle \equiv \frac{1}{T} \int_{t_0}^{t_0+T} Q(t) dt$.

The corresponding *semi-coherent* phase metric \hat{g}_{ij} was first studied in [18] and was found to be expressible as the average over all per-segment coherent phase metrics²) $\tilde{g}_{\ell, ij}$, i.e.

$$\hat{g}_{ij}(\lambda) = \overline{\tilde{g}_{ij}(\lambda)}, \quad (9)$$

where we defined the average operator over segments as

$$\overline{Q} \equiv \frac{1}{N} \sum_{\ell=1}^N Q_{\ell}. \quad (10)$$

The metric is useful for constructing template banks, by providing a simple criterion for how “close” we need to

place templates in order to limit the maximal (relative) loss in detection statistic, which can be written as

$$\max_{\lambda_s} \min_{\lambda_t} \mu(\lambda_s; \lambda_t) \leq \mu_{\text{max}}. \quad (11)$$

This states that the worst-case mismatch to the closest template over the whole template bank should be bounded by a maximal value μ_{max} . Note that each template λ_t “covers” a parameter-space volume given by

$$\mu(\lambda_t; \lambda) = g_{ij}(\lambda_t) \delta\lambda^i \delta\lambda^j \leq \mu_{\text{max}}, \quad (12)$$

which describes an n -dimensional ellipsoid, where n is the number of parameter-space dimensions. The template bank can therefore be thought of as a covering of the whole parameter space with such template ellipsoids, such that no region of \mathcal{P} remains uncovered. One can show [25, 32, 33] that the resulting number of templates \mathcal{N} in such a template bank is expressible as

$$\mathcal{N} = \theta_n \mu_{\text{max}}^{-n/2} \int_{\mathcal{P}} \sqrt{\det g(\lambda)} d^n \lambda, \quad (13)$$

where n is the number of template-bank dimensions, and θ_n is the *center density* (also known as *normalized thickness*) of the covering lattice, which quantifies the number of lattice points per (metric) volume for a unit mismatch. The center density is a geometric property of the lattice, and for the typical covering lattices used here is given by [32]

$$\theta_n = \begin{cases} 2^{-n} n^{n/2} & \text{for } \mathbb{Z}_n, \\ \sqrt{n+1} \left[\frac{n(n+2)}{12(n+1)} \right]^{n/2} & \text{for } A_n^*. \end{cases} \quad (14)$$

An important caveat for using Eq. (13) is that the metric determinant and integration must only include *fully resolved* template-bank dimensions, which require more than one template to cover the extent $\Delta\lambda^i$ of the parameter space along that dimension. We can estimate the per-dimension template extents at given maximal mismatch μ_{max} as³

$$\delta\lambda^i = 2\sqrt{\mu_{\text{max}} [g^{-1}]^{ii}}, \quad (15)$$

where $[g^{-1}]^{ij}$ are the elements of the inverse matrix to the metric g_{ij} . We can therefore define the metric template-bank “thickness” along dimension i in terms of the corresponding effective number of templates along that direction as

$$\mathcal{N}_{\lambda^i} \equiv \frac{\Delta\lambda^i}{\delta\lambda^i}. \quad (16)$$

For dimensions λ^i with $\mathcal{N}_{\lambda^i} \leq 1$ we must exclude this coordinate from the bulk template-counting formula Eq. (13), as it would effectively contribute “fractional templates” and thereby incorrectly *reduce* the total template count.

² This assumes constant signal SNR $\tilde{\rho}_{\ell}(\mathcal{A}, \lambda_s; \lambda_s)$ over all segments ℓ , which should be a good approximation for stationary noise and sufficiently long segments, such that diurnal antenna-pattern variations have averaged out.

³ Note that the extra factor of 2 here compared to [7] is required to account for the total extent of the ellipse, not just the extent from the center. This can also be seen by considering that the center density $\theta_{\mathbb{Z}_1} = 1/2$ in Eq. (13) for the 1-dimensional case.

III. THE BINARY CW PHASE

A. The general phase model

The general CW phase model assumes a slowly spinning-down NS with rotation rate $\nu(\tau)$ and a quadrupolar deformation resulting in the emission of CWs. The phase evolution can therefore be expressed as a Taylor series in the NS source frame as

$$\phi^{\text{src}}(\tau) = 2\pi [f(\tau - t_{\text{ref}}) + \frac{1}{2}\dot{f}(\tau - t_{\text{ref}})^2 + \dots], \quad (17)$$

where t_{ref} denotes the reference time and $f, \dot{f}, \ddot{f}, \dots$ are the CW frequency and spindown parameters. These are defined as

$$f^{(k)} \equiv \sigma \left. \frac{d^k \nu(\tau)}{d\tau^k} \right|_{\tau=t_{\text{ref}}}, \quad (18)$$

where σ is a model-dependent constant relating the instantaneous CW frequency $f(\tau)$ to the NS spin frequency $\nu(\tau)$. For example, for a quadrupolar deformation rotating rigidly with the NS (“mountain”) we have $\sigma = 2$ [2, 9, 34–36], while for r-modes $\sigma \approx 4/3$ [2, 37, 38] and for precession $\sigma \approx 1$ [39, 40].

In order to relate the CW phase in the source frame to the phase $\phi(t_{\text{arr}})$ in the detector frame needed for Eq. (8), we need to relate the wavefront detector arrival time t_{arr} to its source emission time τ , i.e. $\tau(t_{\text{arr}})$, such that $\phi(t_{\text{arr}}) = \phi^{\text{src}}(\tau(t_{\text{arr}}))$. Neglecting relativistic wave-propagation effects⁴ (such as Einstein and Shapiro delays, e.g. see [41, 42] for more details), we can write this as

$$\tau(t_{\text{arr}}) = t_{\text{arr}} + \frac{\vec{r}(t_{\text{arr}}) \cdot \vec{n}}{c} - \frac{d}{c} - \frac{R(\tau)}{c}, \quad (19)$$

where \vec{r} is the vector from solar-system barycenter (SSB) to the detector, c is the speed of light, d is the (generally unknown) distance between the SSB and the binary barycenter (BB), R is the radial distance of the CW-emitting NS from the BB along the line of sight, where $R > 0$ means the NS is further away from us than the BB, and \vec{n} is the unit vector pointing from the SSB to the source. In standard equatorial coordinates with right ascension α and declination δ , the components of the unit vector \vec{n} are given by $\vec{n} = (\cos \alpha \cos \delta, \sin \alpha \cos \delta, \sin \delta)$.

The projected radial distance R along the line of sight can be expressed (e.g. see [43]) as

$$R = y \sin i \sin(\omega + v), \quad (20)$$

where y is the distance of the NS from the BB, i is the inclination angle between the orbital plane and the sky, ω is

the *argument of periaapse*, and v is the *true anomaly* (i.e. the angle from the periaapse to the current NS location around the BB). We further approximate the orbital motion as a pure Keplerian ellipse, which can be described as

$$y(v) = \frac{a(1 - e^2)}{1 + e \cos v}, \quad (21)$$

in terms of the semi-major axis a and the eccentricity e . The ellipse can be written equivalently in terms of the *eccentric anomaly* E , namely

$$y(E) = a(1 - e \cos E), \quad (22)$$

and the dynamics is described by Kepler’s equation, namely

$$\tau - t_p = \frac{P}{2\pi} (E - e \sin E), \quad (23)$$

which provides a (transcendental) relation for $E(\tau)$. Combining Eqs. (20), (21) and (22), we can rewrite the projected radial distance R in terms of E , namely

$$\frac{R}{c} = a_p \left[\sin \omega (\cos E - e) + \cos \omega \sin E \sqrt{1 - e^2} \right], \quad (24)$$

where we defined $a_p \equiv a \sin i / c$. Combining this with Eq. (23) fully determines (albeit only implicitly) the functional relation $R(\tau)$ required for the timing model of Eq. (19).

Dropping the unknown distance d to the BB (which is equivalent to re-defining the intrinsic spindown parameters), and defining the SSB wavefront arrival time t_{SSB} as

$$t_{\text{SSB}}(t_{\text{arr}}; \vec{n}) \equiv t_{\text{arr}} + \frac{\vec{r}(t_{\text{arr}}) \cdot \vec{n}}{c}, \quad (25)$$

we can rewrite the timing relation Eq. (19) as

$$\tau(t_{\text{SSB}}) = t_{\text{SSB}} - \frac{R(\tau)}{c}. \quad (26)$$

Here we are interested only in binary systems with *known* sky-position \vec{n} , and we can therefore change integration variables from t to t_{SSB} in the phase for the metric integration of Eq. (8). Furthermore, given that the Earth’s Rømer delay is bounded by $|\vec{r} \cdot \vec{n} / c| \lesssim 500$ s, and $|dt_{\text{SSB}}/dt| \lesssim 10^{-4}$, this difference will be negligible for the metric and can be dropped. This is equivalent to effectively placing us into the SSB, which is always possible for known \vec{n} . In order to simplify the notation, we now simply write $t \equiv t_{\text{SSB}}$.

Plugging this timing model into the phase of Eq. (17), we obtain⁵

$$\phi(t) \approx 2\pi \left[f \left(\Delta t - \frac{R}{c} \right) + \frac{1}{2} \dot{f} \left(\Delta t - \frac{R}{c} \right)^2 + \dots \right], \quad (27)$$

⁴ These effects are taken into account in the actual matched-filtering search codes, but are negligible for the calculation of the metric.

⁵ The sign on R (and equivalently on a_p) here agrees with Eq.(2.26) in [44], but differs from Eq.(2) in [27], which is incorrect.

with $\Delta t \equiv t - t_{\text{ref}}$. The binary systems we are interested in have semi-major axis a_p of order $\mathcal{O}(1-10)$ s, and binary periods P of order of several hours. Hence, the change in E , and therefore $R(E)$, during the time R/c will be negligible, and so we can approximate $E(\tau) \approx E(t)$, namely

$$t - t_p \approx \frac{P}{2\pi} (E - e \sin E). \quad (28)$$

For the purpose of calculating the metric using Eq. (8), we can further approximate the phase model in the following standard way (e.g. see [28, 30]) as

$$\phi(t) \approx 2\pi \left[(f\Delta t + \frac{1}{2!}\dot{f}\Delta t^2 + \dots) - \frac{R(t)}{c}\bar{f} \right], \quad (29)$$

where we expanded the factors $(\Delta t - R/c)^k$ and kept only the leading-order terms in Δt , keeping in mind that $\Delta t \sim T \gtrsim \mathcal{O}(\text{days})$ and $R/c \sim \mathcal{O}(\text{seconds})$, where T denotes the coherence time. In the last term we replaced the instantaneous intrinsic CW frequency as a function of t by a constant parameter \bar{f} , namely we approximated $\bar{f} \approx (f + \dot{f}\Delta t + \dots)$. This “frequency scale” \bar{f} of the signal could be chosen as the average (or the largest) intrinsic CW frequency of this phase model over the coherence time T . Given that this only enters as a scale parameter in the metric, and the changes of intrinsic frequency over the observation time will generally be small, this introduces only a negligible difference.

B. The small-eccentricity approximation ($e \ll 1$)

We follow the approach of [27] and consider only the *small-eccentricity limit* of the metric. Hence, by Taylor expanding Eqs. (24) and (28) up to leading order in e , i.e. inserting $E(t) = E_0(t) + eE_1(t) + \dots$ into Kepler’s equation Eq. (28), we obtain

$$E_0(t) = \Omega(t - t_p), \quad (30)$$

$$E_1(t) = \sin E_0(t), \quad (31)$$

where $\Omega \equiv \frac{2\pi}{P}$ is the *mean* orbital angular velocity. Plugging this into Eq. (24), we obtain the Rømer delay of the binary to leading order in e as

$$\frac{R}{c} = a_p \left[\sin \psi(t) + \frac{\kappa}{2} \sin 2\psi(t) - \frac{\eta}{2} \cos 2\psi(t) \right], \quad (32)$$

where a constant term $-3a_p\eta/2$ was omitted, which is irrelevant for the metric. We use the standard *Laplace-Lagrange* parameters defined as

$$\kappa \equiv e \cos(\omega), \quad (33)$$

$$\eta \equiv e \sin(\omega), \quad (34)$$

and the mean orbital phase

$$\psi(t) \equiv \Omega(t - t_{\text{asc}}), \quad (35)$$

measured from the *time of ascending node* t_{asc} , which (for small e) is related to t_p by [41]

$$t_{\text{asc}} \equiv t_p - \frac{\omega}{\Omega}, \quad (36)$$

and which (contrary to the time of periape t_p), remains well-defined even in the limit of circular orbits.

The small-eccentricity phase model is therefore parametrized by the 5 binary parameters $\{a_p, t_{\text{asc}}, \Omega, \kappa, \eta\}$ (referred to as the “ELL1” model in [41, 42]), and in the circular case ($e = 0$) this reduces to the 3 binary parameters $\{a_p, t_{\text{asc}}, \Omega\}$ with $\kappa = \eta = 0$.

IV. THE BINARY CW METRIC

A. Phase derivatives

Following [26, 27] we restrict our investigation to (approximately) constant-frequency CW signals, i.e. $f^{(k>0)} = 0$. This is motivated by the assumed steady-state torque-balance situation in LMXBs, which are our main target of interest. However, the corresponding fluctuations in the accretion rate are expected to cause some stochastic frequency drift, and one will therefore need to be careful to restrict the maximal coherence time in order to limit the frequency resolution. This will be discussed in more detail in the application to Scorpius X-1 in Sec. VII. The total phase-evolution parameter space considered here is therefore spanned by the following 6 coordinates

$$\lambda = \{f, a_p, t_{\text{asc}}, \Omega, \kappa, \eta\}. \quad (37)$$

The small-eccentricity phase model used here can now be explicitly written as

$$\frac{\phi(t; \lambda)}{2\pi} \approx f\Delta t - \bar{f}a_p \left[\sin \psi + \frac{\kappa}{2} \sin 2\psi - \frac{\eta}{2} \cos 2\psi \right], \quad (38)$$

with $\psi \equiv \psi(t)$ given by Eq. (35). The frequency parameter \bar{f} is treated as a constant in the phase derivative $\partial_f \phi$ (which corresponds to neglecting the small correction $a_p \ll \Delta t \sim T$), but numerically we have $\bar{f} = f$ in the present constant-frequency case without spindowns. Hence, we obtain the phase derivatives

$$\begin{aligned} \frac{\partial_f \phi}{2\pi} &\approx \Delta t, \\ \frac{\partial_{a_p} \phi}{2\pi} &= -\bar{f} \left[\sin \psi + \frac{\kappa}{2} \sin 2\psi - \frac{\eta}{2} \cos 2\psi \right], \\ \frac{\partial_{t_{\text{asc}}} \phi}{2\pi} &= \bar{f}a_p\Omega [\cos \psi + \kappa \cos 2\psi + \eta \sin 2\psi], \\ \frac{\partial_{\Omega} \phi}{2\pi} &= -\bar{f}a_p(t - t_{\text{asc}}) [\cos \psi + \kappa \cos 2\psi + \eta \sin 2\psi], \\ \frac{\partial_{\kappa} \phi}{2\pi} &= -\frac{1}{2}\bar{f}a_p \sin 2\psi, \\ \frac{\partial_{\eta} \phi}{2\pi} &= \frac{1}{2}\bar{f}a_p \cos 2\psi, \end{aligned} \quad (39)$$

which are inserted into Eq. (8) in order to obtain the coherent metric. The semi-coherent metric is obtained by averaging the coherent metrics over segments according to Eq. (9).

The resulting analytic expressions for these metrics *in the general case* are quite uninformative and unwieldy, while it is straightforward to compute them numerically for any case of interest. However, as noticed in previous investigations [26, 27], it is instructive to focus on two limiting regimes that yield particularly simple analytical results, namely the *long-segment limit* (LS) where $\Delta T \gg P$, and the *short-segment limit* (SS) where $\Delta T \ll P$.

When taking these limits on the metric g_{ij} , in order to decide whether a particular off-diagonal term is negligible or not, it is useful to consider the *diagonal-rescaled* metric

$$g_{ij}^{\text{DN}} \equiv \frac{g_{ij}}{\sqrt{g_{ii} g_{jj}}}, \quad (40)$$

which is dimensionless and has unit diagonal, $g_{ii}^{\text{DN}} = 1$. In this rescaled metric we can then naturally neglect off-diagonal terms if they satisfy $g_{ij}^{\text{DN}} \ll 1$, as their corresponding contribution to the metric mismatch of Eq. (12) will then be negligible.

B. The long-segment (LS) regime ($\Delta T \gg P$)

1. Coherent metric \tilde{g}^{LS}

As noted in [27], it is convenient to use the discrete gauge freedom in the choice of the orbital reference epoch t_{asc} , as we can choose the time of ascending node during any orbit. Therefore we can add any integer multiple p of a period and redefine

$$t'_{\text{asc}} = t_{\text{asc}} + p \frac{2\pi}{\Omega}, \quad p \in \mathbb{Z}, \quad (41)$$

without changing the system. An (infinitesimal) *offset* (or “uncertainty”) δt_{asc} changes under such a transformation in the presence of an (infinitesimal) offset on the period, i.e. if $\delta\Omega \neq 0$, namely

$$\delta t'_{\text{asc}} = \delta t_{\text{asc}} - p \frac{2\pi}{\Omega^2} \delta\Omega. \quad (42)$$

All other coordinate offsets are unaffected by this change of orbital reference epoch. This needs to be carefully taken into account in the metric when redefining t_{asc} . Namely, given that this is a pure “relabeling” of the same physical situation, the corresponding metric mismatch must be invariant, i.e.

$$\begin{aligned} \mu &= \delta\lambda'^m g'_{mb} \delta\lambda'^b \\ &= \delta\lambda^i \left[\frac{\partial\lambda'^m}{\partial\lambda^i} g'_{mb} \frac{\partial\lambda'^b}{\partial\lambda^j} \right] \delta\lambda^j \\ &= \delta\lambda^i g_{ij} \delta\lambda^j, \end{aligned} \quad (43)$$

where the only nonzero components of the Jacobian of this coordinate transformation Eq. (41) are

$$\frac{\partial\lambda'^i}{\partial\lambda^i} = 1, \quad \text{and} \quad \frac{\partial t'_{\text{asc}}}{\partial\Omega} = -p \frac{2\pi}{\Omega^2}. \quad (44)$$

In the long-segment limit, these discrete steps $P = 2\pi/\Omega$ are assumed to be small compared to the segment length ΔT . We can therefore choose the time of ascending node to be approximately at the *segment midtime* t_{mid} , i.e. we consider the special gauge choice $t'_{\text{asc}} \approx t_{\text{mid}}$, in which the phase derivatives Eq. (39) are time-symmetric around t_{mid} , which simplifies the metric calculation. Introducing the time offset

$$\Delta_{\text{ma}} \equiv t_{\text{mid}} - t_{\text{asc}}, \quad (45)$$

we therefor consider first the special case $\tilde{g}_{ij}^{\text{LS}} \equiv \tilde{g}_{ij}^{\text{LS}}(\Delta_{\text{ma}} = 0)$.

Keeping only leading-order terms in $\Omega \Delta T \gg 2\pi$, and up to first order in ϵ (i.e. first order in κ and η), we obtain the diagonal metric

$$\begin{aligned} \tilde{g}_{\text{mid},ii}^{\text{LS}} &= 2\pi^2 \times \\ &\left[\frac{\Delta T^2}{6}, \bar{f}^2, (\bar{f}a_p\Omega)^2, \frac{(\bar{f}a_p\Delta T)^2}{12}, \frac{(\bar{f}a_p)^2}{4}, \frac{(\bar{f}a_p)^2}{4} \right], \end{aligned} \quad (46)$$

in agreement with the result Eq. (30) of [27].

The general case of $\Delta_{\text{ma}} \neq 0$ can be obtained by applying the Jacobian⁶ of Eq. (44) with $\Delta_{\text{ma}} = p 2\pi/\Omega$, and so $t'_{\text{asc}} = t_{\text{asc}} + \Delta_{\text{ma}}$ and $\partial t'_{\text{asc}}/\partial\Omega = -\Delta_{\text{ma}}/\Omega$, resulting in the *general* coherent long-segment metric $\tilde{g}_{ij}^{\text{LS}}(\Delta_{\text{ma}})$ with components

$$\begin{aligned} \tilde{g}_{ff}^{\text{LS}} &= \pi^2 \frac{\Delta T^2}{3}, \\ \tilde{g}_{a_p a_p}^{\text{LS}} &= 2\pi^2 \bar{f}^2, \\ \tilde{g}_{t_{\text{asc}} t_{\text{asc}}}^{\text{LS}} &= 2\pi^2 (\bar{f}a_p\Omega)^2, \\ \tilde{g}_{\Omega\Omega}^{\text{LS}} &= 2\pi^2 (\bar{f}a_p)^2 \left(\frac{\Delta T^2}{12} + \Delta_{\text{ma}}^2 \right), \\ \tilde{g}_{\Omega t_{\text{asc}}}^{\text{LS}} &= \tilde{g}_{t_{\text{asc}}\Omega}^{\text{LS}} = -2\pi^2 (\bar{f}a_p)^2 \Omega \Delta_{\text{ma}}, \\ \tilde{g}_{\kappa\kappa}^{\text{LS}} &= \tilde{g}_{\eta\eta}^{\text{LS}} = \frac{\pi^2}{2} (\bar{f}a_p)^2, \end{aligned} \quad (47)$$

while all other off-diagonal terms are (approximately) zero in this limit. The nonzero off-diagonal term shows that in general (i.e. for $\Delta_{\text{ma}} \neq 0$) there are correlations between offsets in t_{asc} and Ω , introduced by the relation Eq. (42). This generalizes the result in [27] to arbitrary choices of orbital reference epoch t_{asc} .

⁶ A direct but somewhat more complicated calculation of the metric at $\Delta_{\text{ma}} \neq 0$ yields the same result

2. Semi-coherent metric \hat{g}^{LS}

In order to compute the semi-coherent metric, we only need to average the per-segment coherent metrics of the previous sections over segments according to Eq. (9). Here it is crucial to realize that all segment metrics must use the *same coordinates* λ in order for this averaging expression to apply, in particular we can only fix the gauge on t_{asc} *once* for all segments, and so generally $\Delta_{\text{ma},\ell} \neq 0$. This means we *cannot* use the special form of $\hat{g}_{\text{mid}}^{\text{LS}}$ of Eq. (46) for the average over all segments, as done in [27], which would result in the erroneous conclusion that the semi-coherent metric \hat{g}^{LS} would be identical to the coherent one. Instead, using Eq. (9), we find by averaging the per-segment metrics $\hat{g}_\ell^{\text{LS}} \equiv \hat{g}^{\text{LS}}(\Delta_{\text{ma},\ell})$:

$$\begin{aligned}\hat{g}_{ff}^{\text{LS}} &= \pi^2 \frac{\Delta T^2}{3}, \\ \hat{g}_{a_p a_p}^{\text{LS}} &= 2\pi^2 \bar{f}^2, \\ \hat{g}_{t_{\text{asc}} t_{\text{asc}}}^{\text{LS}} &= 2\pi^2 (\bar{f} a_p \Omega)^2, \\ \hat{g}_{\Omega \Omega}^{\text{LS}} &= 2\pi^2 (\bar{f} a_p)^2 \left(\frac{\Delta T^2}{12} + \overline{\Delta_{\text{ma}}^2} \right), \\ \hat{g}_{\Omega t_{\text{asc}}}^{\text{LS}} &= \hat{g}_{t_{\text{asc}} \Omega}^{\text{LS}} = -2\pi^2 (\bar{f} a_p)^2 \Omega \overline{\Delta_{\text{ma}}}, \\ \hat{g}_{\kappa \kappa}^{\text{LS}} &= \hat{g}_{\eta \eta}^{\text{LS}} = \frac{\pi^2}{2} (\bar{f} a_p)^2\end{aligned}\quad (48)$$

It is interesting to consider the ideal case of N regularly-spaced segments without gaps, i.e. $T_{\text{obs}} = N \Delta T$, and

$$\Delta_{\text{ma},\ell} = (\overline{t_{\text{mid}}} - t_{\text{asc}}) + \left(\ell - \frac{N+1}{2} \right) \Delta T, \quad (49)$$

resulting in the first two moments

$$\overline{\Delta_{\text{ma}}} = \overline{t_{\text{mid}}} - t_{\text{asc}}, \quad (50)$$

$$\overline{\Delta_{\text{ma}}^2} = \overline{\Delta_{\text{ma}}}^2 + \frac{1}{12} (N^2 - 1) \Delta T^2, \quad (51)$$

where $\overline{t_{\text{mid}}}$ is the average mid-time over all segments. In this case we can write the variance of $\{\Delta_{\text{ma},\ell}\}$ as

$$\text{var}[\Delta_{\text{ma}}] \equiv \overline{\Delta_{\text{ma}}^2} - \overline{\Delta_{\text{ma}}}^2 = \frac{1}{12} (N^2 - 1) \Delta T^2. \quad (52)$$

If we use the gauge freedom and choose $t_{\text{asc}} \approx \overline{t_{\text{mid}}}$, then $\overline{\Delta_{\text{ma}}} \approx 0$, and $\overline{\Delta_{\text{ma}}^2} \approx (N^2 - 1) \Delta T^2 / 12$, so that in this case the metric is again diagonal, and the only component different from the coherent metric $\hat{g}_{\text{mid}}^{\text{LS}}$ is

$$\hat{g}_{\Omega \Omega}^{\text{LS}} = \pi^2 \frac{(\bar{f} a_p)^2}{6} (N \Delta T)^2 \propto T_{\text{obs}}^2, \quad (53)$$

corresponding to a refinement in the Ω coordinate. Therefore the semi-coherent template-bank spacing in Ω needs to be finer by a factor of N compared to coherent spacing in order to achieve the same mismatch. The existence of this metric refinement in Ω in the long-segment limit had already been noticed earlier [45], and was also used in the discovery of a binary pulsar in Fermi LAT data [46].

C. The short-segment (SS) regime ($\Delta T \ll P$)

1. Coherent metric \hat{g}^{SS}

In the case of very short coherent segments compared to the period, i.e. $\Omega \Delta T \ll 2\pi$, as pointed out in [27], strong parameter-space degeneracies render the metric in these coordinates quite impractical for template-bank generation, and we therefore follow their approach of Taylor-expanding the phase around the mid-point t_{mid} of the coherent observation time, namely

$$\phi(t) = 2\pi \sum_{k=1}^n \frac{u_k}{k!} (t - t_{\text{mid}})^k, \quad (54)$$

omitting a constant term.

The u -coordinates $\{u_k\}$ are defined as the k -th time-derivatives of the phase at t_{mid} , namely

$$u_k \equiv \frac{1}{2\pi} \partial_t^k \phi(t) \big|_{t_{\text{mid}}}. \quad (55)$$

Note that the phase model of Eq. (54) is formally identical to the phase in terms of the frequency and spindowns $\{f^{(k)}\}$ of an isolated NS (with t_{mid} playing the role of the reference time t_{ref}), as can be seen in Eq. (29). The corresponding “spindown metric” has a well-known analytical form, which can be expressed most conveniently in terms of the rescaled dimensionless v -coordinates,

$$v_k \equiv 2\pi \frac{u_k}{k!} \left(\frac{\Delta T}{2} \right)^k, \quad (56)$$

resulting in the v -coordinate metric

$$\begin{aligned}\hat{g}_{kk'}^{\text{SS},v} &= \frac{(k+k'-1) \bmod 2}{k+k'+1} \\ &\quad - \frac{[(k-1) \bmod 2][(k'-1) \bmod 2]}{(k+1)(k'+1)} \\ &= \begin{pmatrix} \frac{1}{3} & 0 & \frac{1}{5} & 0 & \dots \\ 0 & \frac{4}{45} & 0 & \frac{8}{105} & \dots \\ \frac{1}{5} & 0 & \frac{1}{7} & 0 & \dots \\ 0 & \frac{8}{105} & 0 & \frac{16}{225} & \dots \\ \vdots & \vdots & \vdots & \vdots & \ddots \end{pmatrix},\end{aligned}\quad (57)$$

correcting the incorrect expression Eq. (21) of [27].

Applying Eq. (55) to the small-eccentricity phase of Eq. (38), we obtain the following expressions

$$\begin{aligned}u_1 &= -\bar{f} a_p \Omega [\cos \psi_m + \kappa \cos 2\psi_m + \eta \sin 2\psi_m] + f, \\ u_2 &= \bar{f} a_p \Omega^2 [\sin \psi_m + 2\kappa \sin 2\psi_m - 2\eta \cos 2\psi_m], \\ u_3 &= \bar{f} a_p \Omega^3 [\cos \psi_m + 4\kappa \cos 2\psi_m + 4\eta \sin 2\psi_m], \\ u_4 &= -\bar{f} a_p \Omega^4 [\sin \psi_m + 8\kappa \sin 2\psi_m - 8\eta \cos 2\psi_m], \\ u_5 &= -\bar{f} a_p \Omega^5 [\cos \psi_m + 16\kappa \cos 2\psi_m + 16\eta \sin 2\psi_m], \\ u_6 &= \bar{f} a_p \Omega^6 [\sin \psi_m + 32\kappa \sin 2\psi_m - 32\eta \cos 2\psi_m], \\ &\vdots\end{aligned}\quad (58)$$

with $\psi_m \equiv \psi(t_{\text{mid}})$, which differ in the sign of a_p compared to Eqs. (A2-A5) of [27], due to the sign error on R in the phase model discussed in Sec. III A.

Note that for the general elliptical case we expect 6 independent u -coordinates $u(\lambda)$ and 4 in the circular case. The nonlinear coordinate transformation Eqs. (58) can be analytically inverted to obtain $\lambda(u)$, as shown explicitly in Appendix A.

We follow [27] in estimating the range of validity of this Taylor approximation up to order n by considering the order of magnitude of the first neglected term of order $n + 1$, namely

$$|\Delta\phi| \sim 2\pi \frac{u_{n+1}}{(n+1)!} \left(\frac{\Delta T}{2}\right)^{n+1},$$

$$\sim 2\pi \bar{f} a_p \frac{(\pi x)^{n+1}}{(n+1)!}, \quad (59)$$

where we defined the fraction x of an orbit swept during the segment duration, i.e. $x \equiv \Omega\Delta T/(2\pi) = \Delta T/P$. In order to link this phase error to a mismatch, we observe that here $\Delta\phi \sim \Delta v_{n+1}$ of Eq. (56), and the corresponding metric element of Eq. (57) gives $\hat{g}_{(n+1),(n+1)}^{\text{SS},v} \sim \frac{1}{2n+3}$ for $n \gg 1$ and therefore

$$\mu \sim \frac{\Delta\phi^2}{2n+3} = (2\bar{f}a_p)^2 \frac{(\pi x)^{2(n+1)}}{(2n+3)(n+1)!^2}. \quad (60)$$

Plugging in typically values used in the numerical tests later, e.g. $\bar{f} \sim 500\text{Hz}$, $a_p \sim 3\text{s}$, $n = 6$, we see that the mismatch grows very rapidly from $\mu \ll 1$ at $x \lesssim 0.25$ to $\mu \sim 1$ at $x \sim 0.35$. We therefore expect this approximation with $n = 6$ to typically hold for segment durations up to a 1/4 to 1/3 of an orbit.

2. Semi-coherent metric \hat{g}^{SS}

While the segments ΔT are assumed to be short compared to the period P , we only consider the case where the total observation time is long compared to P i.e. in addition to $\Delta T \ll P$ we also assume $T_{\text{obs}} = N\Delta T \gg P$, which also implies $N \gg 1$.

We cannot directly use the coherent per-segment metrics in u -coordinates to compute the semi-coherent metric as an average using Eq. (9), because a fixed physical parameter-space point λ would have different u -coordinates in each segment.

Instead, as first shown in [27], one can go back to physical coordinates and use the fact that $N \gg 1$ to replace the discrete sum in Eq. (9) over segments by an integral over time, namely

$$\hat{g}^{\text{SS}} = \frac{1}{N} \sum_{\ell=1}^N \hat{g}^{\text{SS}}(t_{\text{mid},\ell})$$

$$\stackrel{N \gg 1}{\approx} \frac{1}{T_{\text{obs}}} \int_{t_{\text{mid}} - T_{\text{obs}}/2}^{t_{\text{mid}} + T_{\text{obs}}/2} \hat{g}^{\text{SS}}(t') dt', \quad (61)$$

in terms of coherent per-segment metrics $\hat{g}^{\text{SS}}(t_{\text{mid},\ell})$ expressed as a function of the segment mid-times $t_{\text{mid},\ell}$, and where t_{mid} is the mid-time of the whole observation. In order to compute this, we first calculate the coherent metric of Eq. (8) for each segment $[t_{\text{mid},\ell} - \Delta T/2, t_{\text{mid},\ell} + \Delta T/2]$ by using the phase derivatives of Eq. (39). As before, we keep only first-order terms in e (i.e. first-order in κ, η), then Taylor-expand the results in $\Omega\Delta T \ll 2\pi$ up to second order, obtaining $\hat{g}^{\text{SS}}(t_{\text{mid},\ell})$. This is integrated over the total observation time according to Eq. (61), and only leading-order terms in $\Omega T_{\text{obs}} \gg 2\pi$ are kept. We only keep off-diagonal terms where the corresponding diagonal-rescaled elements Eq. (40) are not $\ll 1$. The resulting semi-coherent short-segment long-observation limit metric elements are found as

$$\hat{g}_{ff}^{\text{SS}} = \pi^2 \frac{\Delta T^2}{3},$$

$$\hat{g}_{a_p a_p}^{\text{SS}} = \frac{\pi^2}{6} (\Omega\Delta T)^2 \bar{f}^2,$$

$$\hat{g}_{t_{\text{asc}} t_{\text{asc}}}^{\text{SS}} = \frac{\pi^2}{6} (\Omega\Delta T)^2 (\bar{f} a_p \Omega)^2, \quad (62)$$

$$\hat{g}_{\Omega\Omega}^{\text{SS}} = \frac{\pi^2}{6} (\Omega\Delta T)^2 (\bar{f} a_p)^2 \left(\frac{T_{\text{obs}}^2}{12} + \overline{\Delta_{\text{ma}}}^2 \right),$$

$$\hat{g}_{\Omega t_{\text{asc}}}^{\text{SS}} = \hat{g}_{t_{\text{asc}} \Omega}^{\text{SS}} = -\frac{\pi^2}{6} (\Omega\Delta T)^2 (\bar{f} a_p)^2 \Omega \overline{\Delta_{\text{ma}}},$$

$$\hat{g}_{\kappa\kappa}^{\text{SS}} = \hat{g}_{\eta\eta}^{\text{SS}} = \frac{\pi^2}{6} (\Omega\Delta T)^2 (\bar{f} a_p)^2,$$

which generalizes the result in [27] (for gauge choice $t_{\text{asc}} \approx \bar{t}_{\text{mid}}$, i.e. $\overline{\Delta_{\text{ma}}} \approx 0$) to the general case of $\overline{\Delta_{\text{ma}}} \neq 0$.

Note that a more general form encompassing the semi-coherent metric in both the short-segment and long-segment limits has recently found in [47].

V. NUMBER OF TEMPLATES

In order to express the explicit template counts based on Eq. (13), we will assume a simple parameter space \mathcal{P} bounded by $\bar{f} \in [\bar{f}_{\text{min}}, \bar{f}_{\text{max}}]$, $a_p \in [a_{p\text{min}}, a_{p\text{max}}]$, $t_{\text{asc}} \in [t_{\text{ascmin}}, t_{\text{ascmax}}]$, $\Omega \in [\Omega_{\text{min}}, \Omega_{\text{max}}]$, $e \in [e_{\text{min}}, e_{\text{max}}]$, and $\omega \in [\omega_{\text{min}}, \omega_{\text{max}}]$. Using Eq. (33), (34) we see that

$$\int \int d\kappa d\eta = \int \int e de d\omega. \quad (63)$$

Observing that $g_{\kappa\kappa} = g_{\eta\eta}$, we can further obtain $\mu = g_{\kappa\kappa} d\kappa^2 + g_{\eta\eta} d\eta^2 = g_{\kappa\kappa} (de^2 + e^2 d\omega^2)$, which implies

$$g_{ee} = g_{\kappa\kappa}, \quad g_{\omega\omega} = e^2 g_{\kappa\kappa}, \quad g_{e\omega} = 0. \quad (64)$$

For convenience of notation, for the following template-count expressions, we define the shorthand

$$|Q| \equiv Q_{\text{max}} - Q_{\text{min}}. \quad (65)$$

A. Templates \mathcal{N}^{LS} in the long-segment regime

The determinant of the semi-coherent metric \hat{g}^{LS} of Eq. (48) only differs from that of the coherent metric \tilde{g}^{LS} of Eq. (47) by the presence of a refinement factor γ , i.e.

$$\det \hat{g}^{\text{LS}} = \gamma^2 \det \tilde{g}^{\text{LS}}, \quad \text{with} \quad (66)$$

$$\gamma \equiv \sqrt{1 + 12 \frac{\text{var}[\Delta_{\text{ma}}]}{\Delta T^2}} \stackrel{\text{gapless}}{=} N, \quad (67)$$

where the last equality holds in the special case of N segments without gaps, as seen from Eq. (52). This refinement only affects the Ω -dimension. In particular, for the gauge choice $\Delta_{\text{ma}} = 0$, we see

$$\hat{g}_{\text{mid},\Omega}^{\text{LS}} = \gamma^2 \tilde{g}_{\text{mid},\Omega}^{\text{LS}}. \quad (68)$$

The determinant Eq. (66) is seen to be independent of the gauge choice on t_{asc} (i.e. on Δ_{ma}), and we can therefore write the template volume density in factored form as

$$\sqrt{\det g^{\text{LS}}} d^n \lambda = \prod_{i=1}^n \sqrt{g_{\text{mid},ii}^{\text{LS}}} d\lambda^i. \quad (69)$$

This independence of Δ_{ma} in the volume element may seem surprising at first, as one would expect [48] the parameter-space uncertainty on t_{asc} to grow with increasing separation from the observation time $t_{\text{mid}} = t_{\text{asc}} + \Delta_{\text{ma}}$, if there is any uncertainty $d\Omega$ in the mean angular velocity, as also discussed in Sec. IV B. However, this is only the case if the uncertainty in Ω is not resolved by the template bank, when one observes indeed a stretching of the effective (projected) uncertainty in t_{asc} , resulting in an increase in required templates $\mathcal{N}_{t_{\text{asc}}}$. On the other hand, if the $\Omega - t_{\text{asc}}$ parameter space is fully resolved by the metric [as assumed in Eq. (69)], there is no increase of templates. Changing Δ_{ma} only deforms the $\Omega - t_{\text{asc}}$ parameter space in a volume-preserving way by shear according to Eq. (41). Given that any offset in t_{asc} and Ω results in a deterministically shifted offset t'_{asc} , there is no loss of information, and therefore no increase in the number of templates. Alternatively, one can consider the orbital epoch t_{asc} with its uncertainty fixed at its original measurement epoch, and account for the offset Δ_{ma} in the metric as done in Sec. IV B. This results in changing metric correlations between t_{asc} and Ω while leaving the ellipse volume unchanged.

The expression for the number of templates per dimension \mathcal{N}_{λ^i} of Eqs. (15) and (16) simplifies to

$$\mathcal{N}_{\lambda^i}^{\text{LS}} = \frac{1}{2} \mu_{\text{max}}^{-1/2} \sqrt{g_{\text{mid},ii}^{\text{LS}}} \Delta \lambda^i, \quad (70)$$

with individual components in coordinates $\lambda^i = \{f, a_p, t_{\text{asc}}, \Omega, e\}$ given by

$$\sqrt{g_{\text{mid},ii}^{\text{LS}}} = \pi \sqrt{2} \left[\frac{\Delta T}{\sqrt{6}}, \bar{f}, \bar{f} a_p \Omega, \bar{f} a_p \frac{\gamma \Delta T}{\sqrt{12}}, \frac{\bar{f} a_p}{2} \right], \quad (71)$$

noting that $g_{ee} = g_{\kappa\kappa} = g_{\eta\eta}$ and $g_{\omega\omega} = e^2 g_{ee}$. In practice, one often encounters cases where the parameter uncertainty along some of these dimensions is smaller than the metric resolution, such that a single template covers the whole extent of the parameter space along that direction. In this case the corresponding coordinate contribution to the template density would result in fractional templates (and therefore underestimating the number of templates) and must not be included in Eq. (13), as discussed in Sec. II C.

The number of templates over the full 6D parameter space $\{f, a_p, t_{\text{asc}}, \Omega, e, \omega\}$ according to Eq. (13) is obtained as

$$\mathcal{N}^{\text{LS}} = \frac{\theta_6}{\mu_{\text{max}}^3} \frac{\pi^6 \gamma \Delta T^2}{360 \sqrt{2}} [\bar{f}^6] [a_p^5] [t_{\text{asc}}] [\Omega^2] [e^2] [\omega]. \quad (72)$$

In the application to Scorpius X-1, a few special cases will be of interest, namely when one or more of the uncertainties in orbital parameters are smaller than the template extent.

For sufficiently well-estimated orbital angular velocity $\Omega = \Omega_0$ (such that $\mathcal{N}_\Omega < 1$), the template count for the corresponding 5D template bank over $\{f, a_p, t_{\text{asc}}, e, \omega\}$ is found as

$$\mathcal{N}^{\text{LS}} = \frac{\theta_5}{\mu_{\text{max}}^{5/2}} \frac{\pi^5 \Delta T}{40 \sqrt{3}} \Omega_0 [\bar{f}^5] [a_p^4] [t_{\text{asc}}] [e^2] [\omega], \quad (73)$$

while in the 4D case of well-estimated e and ω (or equivalently, for circular orbits) $\{f, a_p, t_{\text{asc}}, \Omega\}$, we obtain

$$\mathcal{N}^{\text{LS}} = \frac{\theta_4}{\mu_{\text{max}}^{5/2}} \frac{\pi^4 \gamma \Delta T^2}{36 \sqrt{2}} [\bar{f}^4] [a_p^3] [t_{\text{asc}}] [\Omega^2], \quad (74)$$

and finally in the 3D circular case with well-determined $\Omega = \Omega_0$, e and ω we find the template count over the remaining parameter space $\{f, a_p, t_{\text{asc}}\}$ as

$$\mathcal{N}^{\text{LS}} = \frac{\theta_3}{\mu_{\text{max}}^{3/2}} \frac{\pi^3 \Delta T}{\sqrt{27}} \Omega_0 [\bar{f}^3] [a_p^2] [t_{\text{asc}}]. \quad (75)$$

We note that Eqs. (71)–(75) are valid for the semi-coherent case with general refinement γ , as well as for the coherent case with $\gamma = 1$.

B. Templates \mathcal{N}^{SS} in the short-segment regime

Due to the nonlinear transformation Eqs. (58) from physical parameters λ into u -coordinates, the physical parameter space \mathcal{P} would be described by complicated integration boundaries in u . In addition, the number of dimensions in u -coordinates that need to be included in the template bank is a non-trivial function of the signal parameters and the mismatch μ_{max} (e.g. as seen later in Fig. 3). These effects would dominate the expression for the number of templates Eq. (13), and it is therefore not clear whether a useful closed analytical form can be given.

Using the metric volume density for the semi-coherent short-segment metric \hat{g}^{SS} of Eq. (62), we obtain the following number of templates Eq. (13) for the full 6D template bank

$$\hat{N}^{\text{SS}} = \frac{\theta_6}{\mu_{\text{max}}^3} \frac{\pi^6 \Delta T^6 T_{\text{obs}}}{90720\sqrt{6}} [\bar{f}^6] [a_{\text{p}}^5] [t_{\text{asc}}] [\Omega^7] [e^2] [\omega], \quad (76)$$

while in the 4D circular case we find

$$\hat{N}^{\text{SS}} = \frac{\theta_4}{\mu_{\text{max}}^2} \frac{\pi^4 \Delta T^4 T_{\text{obs}}}{2160\sqrt{6}} [\bar{f}^4] [a_{\text{p}}^3] [t_{\text{asc}}] [\Omega^5]. \quad (77)$$

VI. NUMERICAL TESTS OF THE METRICS

In this section we present numerical tests performed on the parameter-space metrics derived in the previous sections. These tests consist in comparing the predicted metric mismatches μ of Eq. (7) against measured \mathcal{F} -statistic mismatches μ_0 of Eq. (5) obtained via signal software injections. The quantity used for this comparison is the (symmetric) relative difference [29–31], defined as

$$\varepsilon(\mu_0, \mu) \equiv \frac{\mu_0 - \mu}{0.5(\mu_0 + \mu)}, \quad (78)$$

which is bounded within $\varepsilon \in [-2, 2]$. For small values $|\varepsilon| \ll 1$ this agrees with the asymmetric definitions of relative errors, such as $\varepsilon_0 \equiv (\mu_0 - \mu)/\mu_0$, with domain $\varepsilon_0 \in (-\infty, 1]$, and which is related to the symmetric relative difference ε as $\varepsilon_0 = \frac{\varepsilon}{1+\varepsilon/2}$.

A. Monte-Carlo software-injection method

The general algorithm used for these software-injection tests is the following:

1. Pick random signal amplitude parameters \mathcal{A} and phase-evolution parameters λ_{s} from suitable priors.
2. Generate a phase-parameter offset $\delta\lambda = \lambda_{\text{t}} - \lambda_{\text{s}}$ by finding the closest lattice template λ_{t} to the signal point λ_{s} .
3. Compute the phase metric $g_{ij}(\lambda_{\text{s}})$ and metric mismatch $\mu(\lambda_{\text{s}}, \lambda_{\text{t}}) = g_{ij} \delta\lambda^i \delta\lambda^j$ according to Eq. (7).
4. Generate a (noise-free) data-set containing the signal $x = h(t; \mathcal{A}, \lambda_{\text{s}})$. Compute the (coherent or semi-coherent) \mathcal{F} -statistic at the injection point, $\mathcal{F}(x; \lambda_{\text{s}})$, and at the template, $\mathcal{F}(x; \lambda_{\text{t}})$. In noise-free data we have $\mathcal{F}(x; \lambda) = E[\mathcal{F}(x; \lambda)]$, therefore we can obtain $\rho^2(\mathcal{A}, \lambda_{\text{s}}; \lambda)$ via Eq. (1) or Eq. (4), respectively, which yields the measured mismatch μ_0 via Eq. (5). Injection and recovery is performed using the LALSuite software package [49].

This procedure is applied to test each of the four types of metric detailed in Secs. IV B and IV C.

Choice of signal parameters $\{\mathcal{A}, \lambda_{\text{s}}\}$

The random signal amplitude parameters \mathcal{A} are chosen as follows: the scalar amplitude h_0 plays no role for the metric and is fixed to $h_0 = 1$. The inclination angle is drawn from a uniform distribution in $\cos \iota \in [-1, 1]$, the polarization angle within $\psi \in [0, 2\pi]$, and the (irrelevant) initial phase within $\phi_0 \in [0, 2\pi]$.

The sky-position for all signals is fixed (without loss of generality) to that of Scorpius X-1, namely $(\alpha, \delta) = (4.276, -0.273)$ rad, and we assume gapless data from the LIGO Hanford detector (H) [3].

The random signal phase-evolution parameters $\lambda_{\text{s}} \in \mathcal{P}$ are generated by drawing them from uniform distributions over the ranges:

$$\begin{aligned} f &\in [50, 1000] \text{ Hz}, \\ a_{\text{p}} &\equiv \frac{a \sin i}{c} \in [1, 5] \text{ s}, \\ P &\in [P_0 - dP, P_0 + dP], \\ t_{\text{p}} &\in \left[t_{\text{mid}} - \frac{P}{2}, t_{\text{mid}} + \frac{P}{2} \right], \\ \log_{10} e &\in [-5, \log_{10}(0.9)], \\ \omega &\in [0, 2\pi] \text{ rad}. \end{aligned} \quad (79)$$

The orbital period “scale” P_0 is fixed for each set of software injections (specified later), while the corresponding sampling range $dP = P_0^2 d\Omega/(2\pi)$ is given in terms of a range $d\Omega$ in orbital angular velocity, which is chosen as $d\Omega \sim 1/\sqrt{g\Omega}$, corresponding to a roughly unity mismatch along Ω . The motivation for this construction is twofold: on one hand we need to control the scale of the orbital period in order to ensure the appropriate short-segment ($\Delta T \ll P$) or long-segment ($\Delta T \gg P$) limit is satisfied in the tests. On the other hand, we wish to randomize the period over a range larger than the typical template-bank spacings in order to fully sample the Wigner-Seitz cell of the template bank.

For the metric tests presented in this section we employ a Demodulation method of computing the \mathcal{F} -statistic, using Short Fourier Transforms (SFTs) of length $T_{\text{sft}} = 5$ s. Due to the linear phase model over each SFT employed by the search code, this corresponds to a maximum error in phase of $|\Delta\phi| \sim 0.07$ rad over the investigated parameter space (see Appendix C for more details).

Generating template-bank offsets $\delta\lambda$

Sampling suitable offsets $\delta\lambda$ for metric testing has proved to be a subtle and difficult point in previous metric studies [29, 30]. Here we employ a recent innovation for generating “natural” offsets from a *virtual* template bank (e.g. see [31] for more details). This method can be applied whenever the metric g_{ij} is constant (i.e. independent of λ), which allows for constructing a lattice-based

template bank for a given maximal mismatch μ_{\max} of Eq. (12), namely

$$\mu_{\max} \geq \delta \lambda^i g_{ij} \delta \lambda^j = \delta \beta^l \delta_{lk} \delta \beta^k, \quad (80)$$

where $\delta \beta^l \equiv \delta \lambda^i A_i^l$ in terms of the Cholesky decomposition of the metric, i.e. $g_{ij} = A_i^l \delta_{lk} A_j^k$. Further rescaling $\delta q^i \equiv (\mathcal{R}/\sqrt{\mu_{\max}}) \delta \beta^i$ yields the corresponding Euclidean covering equation

$$\delta q^i \delta_{ij} \delta q^j \leq \mathcal{R}^2, \quad (81)$$

in terms of the covering radius \mathcal{R} . This requirement can be satisfied by lattices [32, 50] such as, for example, the simple n -dimensional hyper-cubic lattice \mathbb{Z}_n (with covering radius of $\mathcal{R}_{\mathbb{Z}_n} = \sqrt{n}/2$), or the highly efficient covering lattice A_n^* . For both \mathbb{Z}_n and A_n^* lattices, efficient algorithms exist for finding the closest lattice point q_t , satisfying Eq. (81) for any given point $q_s \in \mathbb{R}^n$, with $\delta q = q_t - q_s$. For example, for \mathbb{Z}_n this is trivially given by component-wise rounding, i.e. $q_t = \text{round}[q_s]$. We can use this construction to: (i) easily find the closest lattice template λ_t to any given λ_s , (ii) transform into lattice coordinates, i.e. $q_s = q(\lambda_s)$, (iii) find the closest lattice point q_t and invert back, i.e. $\lambda_t = \lambda(q_t)$. The resulting distribution of offsets $\delta \lambda = \lambda_t - \lambda_s$ uniformly samples the Wigner-Seitz cell of the corresponding lattice, which is the natural offset distribution for uniform signal probability over the parameter space $\lambda_s \in \mathcal{P}$.

Note that while the coherent short-segment metric \tilde{g}^{SS} of Eq. (57) is *explicitly flat* (i.e. constant), this is not true for the other three cases, namely the semi-coherent short-segment metric of Eq. (62), the coherent long-segment metric of Eq. (47), and the semi-coherent long-segment metric of (48). However, it is easy to see that diagonal-rescaling via Eq. (40) achieves explicit flatness in all these three cases, namely by working in terms of rescaled coordinates

$$\delta \lambda'^i \equiv \sqrt{g_{ii}} \delta \lambda^i. \quad (82)$$

It might seem that the terms involving $\Delta_{\text{ma}} \equiv t_{\text{mid}} - t_{\text{asc}}$ depend on the coordinate t_{asc} , but this effect can be neglected in all these cases, as we always assume $T_{\text{obs}} \gg P$, and by choosing a gauge on t_p as in Eq. (79), we can effectively approximate $\Delta_{\text{ma}}/T_{\text{obs}} \approx 0$ in the metric.

The following metric tests use a \mathbb{Z}_6 lattice and a maximum mismatch of $\hat{\mu}_{\max} = \tilde{\mu}_{\max} = 0.3$. This lattice is chosen purely for simplicity, and only serves to generate realistic signal offsets for testing the metric. Using a different lattice (e.g. A_n^*) would slightly change the distribution of sampled offsets, but would be inconsequential for the metric-testing results. The maximal-mismatch value of $\mu_{\max} = 0.3$ represents a fairly “typical” value for realistic searches, but is still small enough so that measured \mathcal{F} -statistic losses are only minimally affected by higher-order corrections compared to the metric approximation (e.g. see [29, 30]). Here we are mostly interested in the accuracy of the metric within its range of applicability, while nonlinear deviations from the metric approximation would warrant a separate study.

B. Results in the long-segment regime

1. Coherent long-segment metric \tilde{g}^{LS}

For these tests we fix the “scale” of the orbital period to $P_0 = 19$ h (similar to Scorpius X-1), and vary the segment length ΔT from 3 days to 39 days in steps of 4 days. We choose the parameter space range for the orbital velocity Ω as two times the largest possible metric spacing occurring over the parameter space, i.e., $d\Omega = 2 \max[(\tilde{g}_{\Omega\Omega}^{\text{LS}})^{-1/2}] \sim 1.2 \times 10^{-7} \text{ s}^{-1}$. The resulting parameter-space half-width in Period is $dP \sim 90$ s, and therefore we have at least $\Delta T/P \geq 3.8$.

The results of these tests are shown in Fig. 1 for the total 20 000 trials performed (2 000 trials for each ΔT value). We note that the performance of the parameter-space metric starts to degrade below $\Delta T/P \lesssim 10$ for low-eccentricity orbits [panel (b)], and becomes generally poor for high-eccentricity orbits above $e \gtrsim 10^{-1}$ [panels (a) and (c)]. Panels (d) and (e) show the agreement between the measured mismatch distribution and the expected \mathbb{Z}_6 -lattice distribution. In panel (f) we see that the phase-metric approximation of Eq. (6) agrees better for small mismatch values and develops a slight tendency to *over-estimate* the actual loss $\tilde{\mu}_0$ for higher mismatches. This is a general feature seen in all four cases tested here, and qualitatively agrees with a similar effect seen in tests of the all-sky metric for isolated CW signals [29, 30].

2. Semi-coherent long-segment metric \hat{g}^{LS}

The scale for the orbital period used here is $P_0 = 2$ h, with segments of fixed length of $\Delta T = 1$ day and varying T_{obs} . In total we used 10 values of T_{obs} : 1 day, 10 days, and 30 days up to 100 days in steps of 10 days.

The parameter-space extent in orbital velocity is the same as considered in the previous section, i.e., $d\Omega \sim 1.2 \times 10^{-7} \text{ s}^{-1}$, resulting in a parameters-space half-width for the period of $dP \sim 1$ s. Therefore $\Delta T/P \geq 12$ is satisfied for all trials.

The results of these tests are shown in Fig. 2 for the total of 20 000 trials performed (2 000 trials for each T_{obs} value). From panels (a, b, c) we see that the agreement between measurements and predictions is generally good in the low-eccentricity regime $e < 10^{-1}$, while rapidly degrading at higher eccentricities. The case $T_{\text{obs}} = \Delta T = 1$ day corresponds to the single-segment coherent case, and antenna-pattern effects are expected to play a role for observation times of order a day.

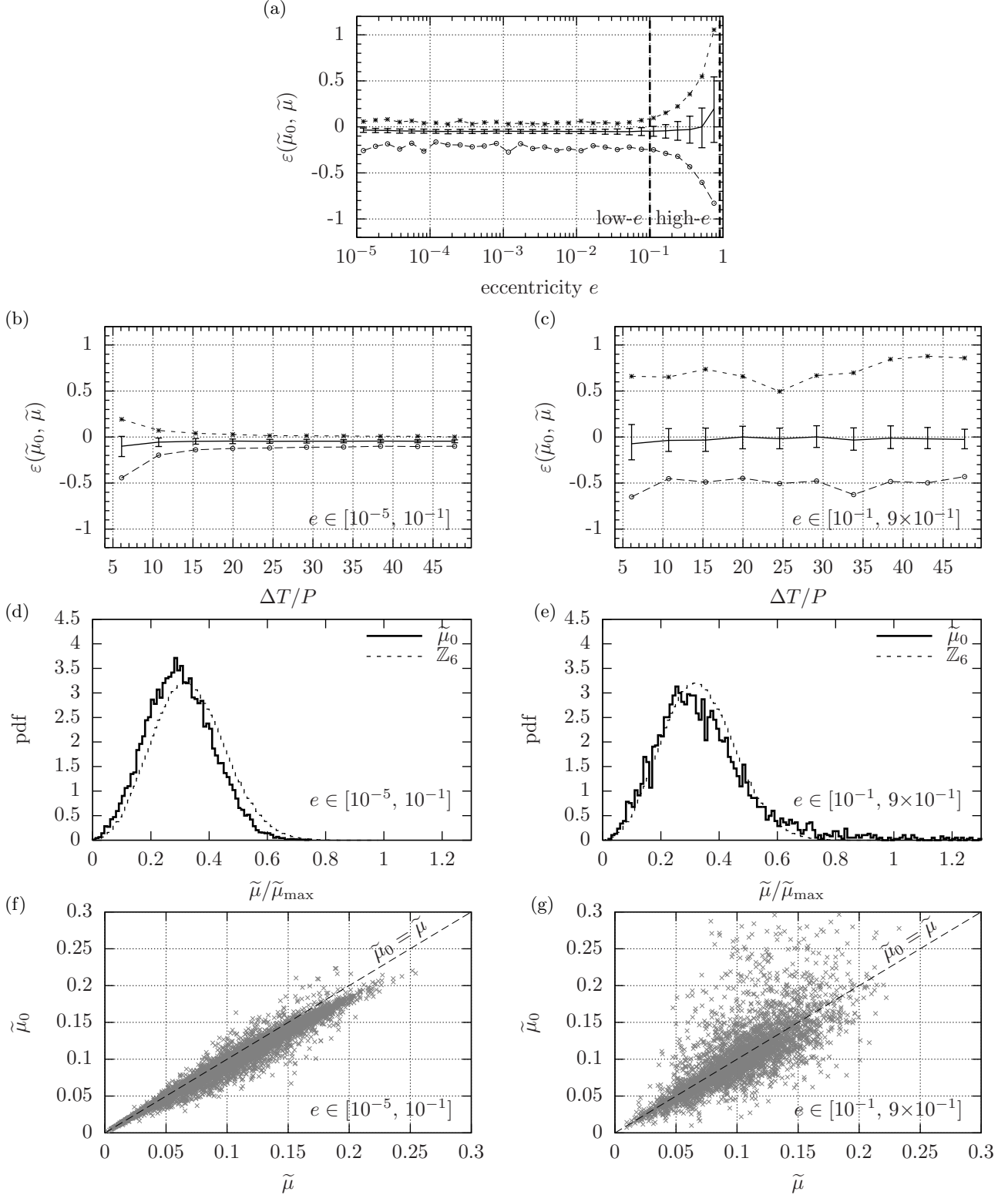


FIG. 1. *Coherent long-segment regime*: results for metric tests on \tilde{g}^{LS} of Sec. IV B 1, for orbital period scale of $P_0 = 19$ h and varying segment length ΔT from 3 to 39 days in steps of 4 days. (a) Relative error $\varepsilon(\tilde{\mu}_0, \tilde{\mu})$ versus eccentricity e . The dashed vertical line denotes the boundary between a “low- e ” range $e \in [10^{-5}, 0.1]$ (left-column plots) and a “high- e ” range $e \in [0.1, 0.9]$ (right-column plots). (b,c): Relative error ε vs observation time $\Delta T/P$. The solid lines in panels (a, b, c) denote the median value, the error bars correspond to the 25th-75th percentiles, the circles and the stars denote the 2.5th and 97.5th percentiles, respectively. (d,e): Mismatch histogram of measured mismatches $\tilde{\mu}_0/\tilde{\mu}_{\max}$ and theoretical distribution in a \mathbb{Z}_6 lattice. (f,g): Measured mismatch $\tilde{\mu}_0$ versus predicted $\tilde{\mu}$.

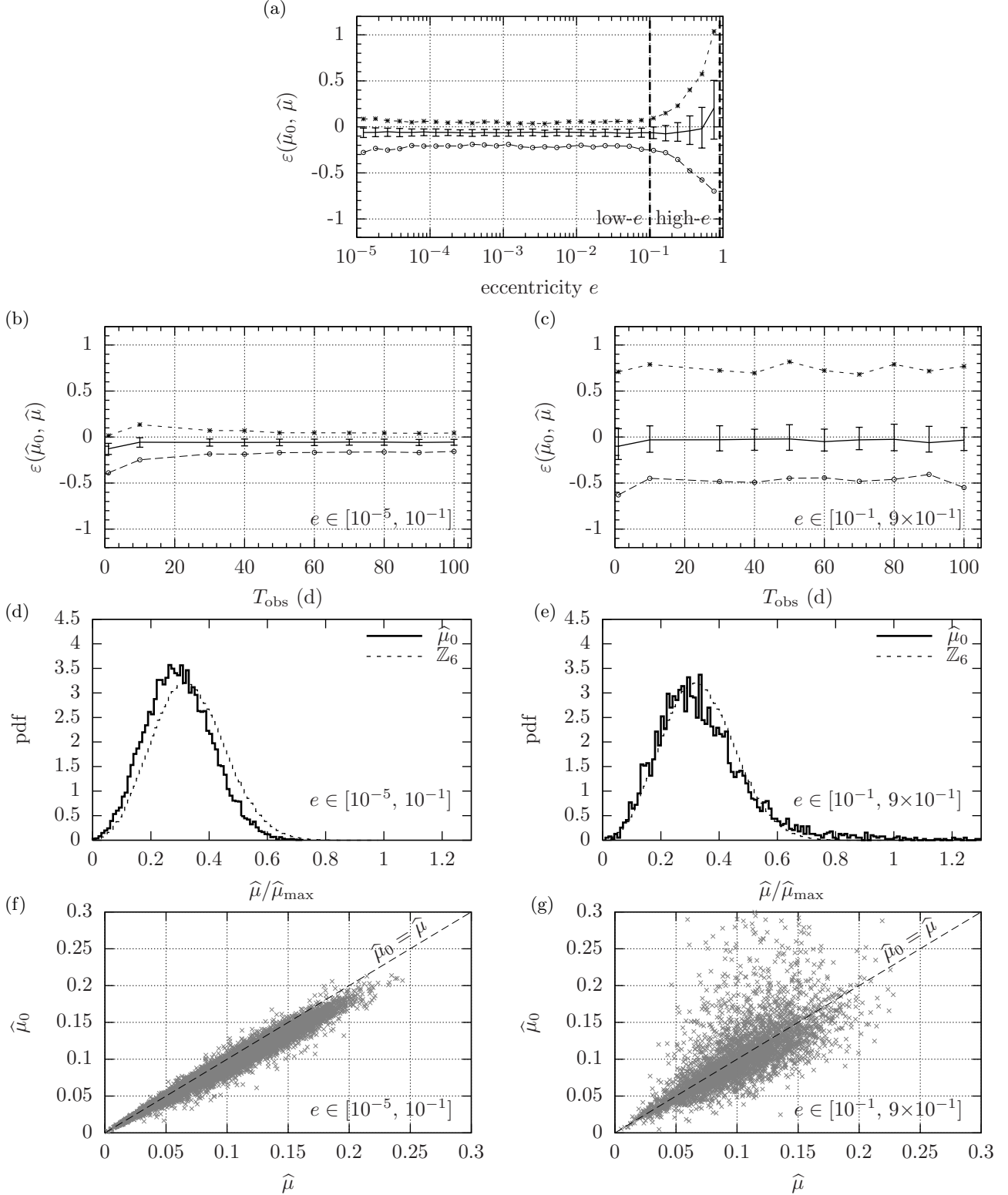


FIG. 2. *Semi-coherent long-segment* regime: results for metric tests on \hat{g}^{LS} of Sec. IV B 2 for orbital period scale of $P_0 = 2$ h and fixed segment length of $\Delta T = 1$ day. (a) Relative error $\varepsilon(\hat{\mu}_0, \hat{\mu})$ versus eccentricity e . The dashed vertical line denotes the boundary between a “low- e ” range $e \in [10^{-5}, 0.1]$ (left-column plots) and a “high- e ” range $e \in [0.1, 0.9]$ (right-column plots). (b,c): Relative error ε vs observation time T_{obs} . The solid lines in panels (a, b, c) denote the median value, the error bars correspond to the 25th-75th percentiles, the circles and the stars denote the 2.5th and 97.5th percentiles, respectively. (d,e): Mismatch histogram of measured mismatches $\hat{\mu}_0/\hat{\mu}_{\text{max}}$ and theoretical distribution in a \mathbb{Z}_6 lattice. (f,g): Measured mismatch $\hat{\mu}_0$ versus predicted $\hat{\mu}$.

C. Results in the short-segment regime

1. Coherent short-segment metric \tilde{g}^{SS}

There is an important technical difference in the injection algorithm in this case, as the metric is expressed in u -coordinates of Eq. (58) instead of physical coordinates λ . Here the physical coordinates of an injected signal λ_s are converted into u -coordinates, $u_{s,k} \equiv u_k(\lambda_s)$, and then the closest lattice template $u_{t,k}$ is found. In principle one could try to convert this back into physical coordinates using the expressions given in Appendix A, but this is not possible in all cases, as some templates in u -coordinates do not correspond to physical coordinates and therefore cannot be inverted. In order to circumvent this problem, we use the fact that the u -coordinates essentially describe an isolated CW signal with $n - 1 = 5$ spindown values. Hence, the perfectly-matched \mathcal{F} -statistic is computed as usual at $\mathcal{F}_s \equiv \mathcal{F}(\lambda_s)$, but the \mathcal{F} -statistic for the mismatched template is computed at the isolated spindown location defined by the u -coordinates, namely by setting $f_t^{(k)} = u_{t,k+1}$ for $k = 0, \dots, 5$, i.e. $\mathcal{F}_t \equiv \mathcal{F}(f_t^{(k)})$.

The orbital period scale is set as $P_0 = 80$ days, and the resolution for the orbital angular frequency is taken to be $d\Omega = 2 \times 10^{-7} \text{ s}^{-1}$, and hence $P \in [62, 97]$ days. We sample segment lengths ΔT ranging from 2 days up to 25 days in steps of half a day, exploring the range in relative segment length of $\Delta T/P \sim [0.02, 0.4]$, using a total of 94 000 trials performed (2 000 trials for each ΔT value).

In order to better understand these results, it is interesting to consider the effective template dimension for the different software injections, which we can quantify in terms of the highest nonzero component $u_{t,k'}$ of the closest template found for an injection, namely

$$k' \equiv \max k \text{ with } |u_{t,k}| > 0. \quad (83)$$

This quantity is plotted Fig. 3 for the injections per-

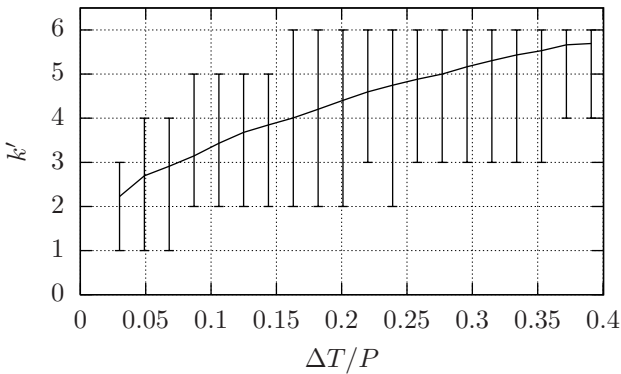


FIG. 3. Effective template dimension k' of Eq. (83) versus segment length $\Delta T/P$ for the software-injection studies shown in Fig. 4. Solid line: mean value, error bars: maximum and minimum of k' .

formed for Fig. 4, giving the range of effective template dimensions k' as a function of relative segment length $\Delta T/P$. As expected, we see that the effective template dimension increases with $\Delta T/P$ and eventually would require more than 6 u -coordinates for $\Delta T/P \gtrsim 0.2$, corresponding to the region where the u -metric performance is seen in Fig. 4(b) to start to degrade. We also see that the average template-bank dimension is not constant and generally less than 6, which explains the discrepancy between the sampled mismatch distribution in Fig. 4(d) compared to the expected \mathbb{Z}_6 -lattice distribution.

The results of the injection tests are shown in Fig. 4. From panels (b) and (c) we see that the approximation in terms of 6 u -coordinates used here starts to break down in the range $\Delta T/P \gtrsim [0.25, 0.35]$, as anticipated from Eq. (60) and Fig. 3. Also, contrary to the other three limits tested, the small-eccentricity approximation only seems to work well for eccentricities up to about $e \lesssim 3 \times 10^{-3}$, after which it rapidly deteriorates and increasingly loses predictive power beyond $e \gtrsim 10^{-2}$ [see panel (a)].

2. Semi-coherent short-segment metric \tilde{g}^{SS}

The scale for the orbital period here is fixed to $P_0 = 10$ days, and we used a fixed segment length of $\Delta T = 1$ day. The parameter space extent in terms of Ω is $d\Omega = 2 \max[(\tilde{g}_{\Omega\Omega}^{\text{SS}})^{-1/2}] \sim 6.6 \times 10^{-8} \text{ s}^{-1}$ resulting in $dP \sim 2.2$ hours and therefore we have at most $\Delta T/P \lesssim 0.12$. We use a varying T_{obs} ranging from 30 up to 100 days in steps of 10 days. The results of these tests are shown in Fig. 5 for the total 16 000 trials performed (2 000 trials for each T_{obs} value). We note that the validity extent of the semi-coherent metric is constantly acceptable for low-eccentricity orbits [panel (b)] and degrades (albeit much less than in the other three cases) for higher eccentricity $e \gtrsim 10^{-1}$ [panels (a) and (c)]. Panels (d) and (e) compare the mismatch histogram of measured mismatches with the theoretical distribution in a \mathbb{Z}_6 lattice. Furthermore, we also note that in this regime the metric mismatch $\hat{\mu}$ tends to generally *over-estimate* $\hat{\mu}_0$ somewhat more strongly than in the other cases [see panels (f) and (g)].

D. General discussion of metric testing results

Summarizing the results presented in Figs. 1- 5, we observe that the agreement between metric mismatch predictions and the measured relative loss in \mathcal{F} -statistic is generally very good (typically no worse than ~ 10 – 30%) within its range of applicability. In the following we further quantify where various approximations start to break down.

The first-order small-eccentricity approximation (see Sec. III B) starts to noticeably degrade only above $e \gtrsim 0.1$ [Fig. 1(a), Fig. 2(a) and Fig. 5(a)] except in the coherent short-segment regime where it breaks down above $e \gtrsim 5 \times 10^{-3}$ [Fig. 4(a)]. Further investigation of this

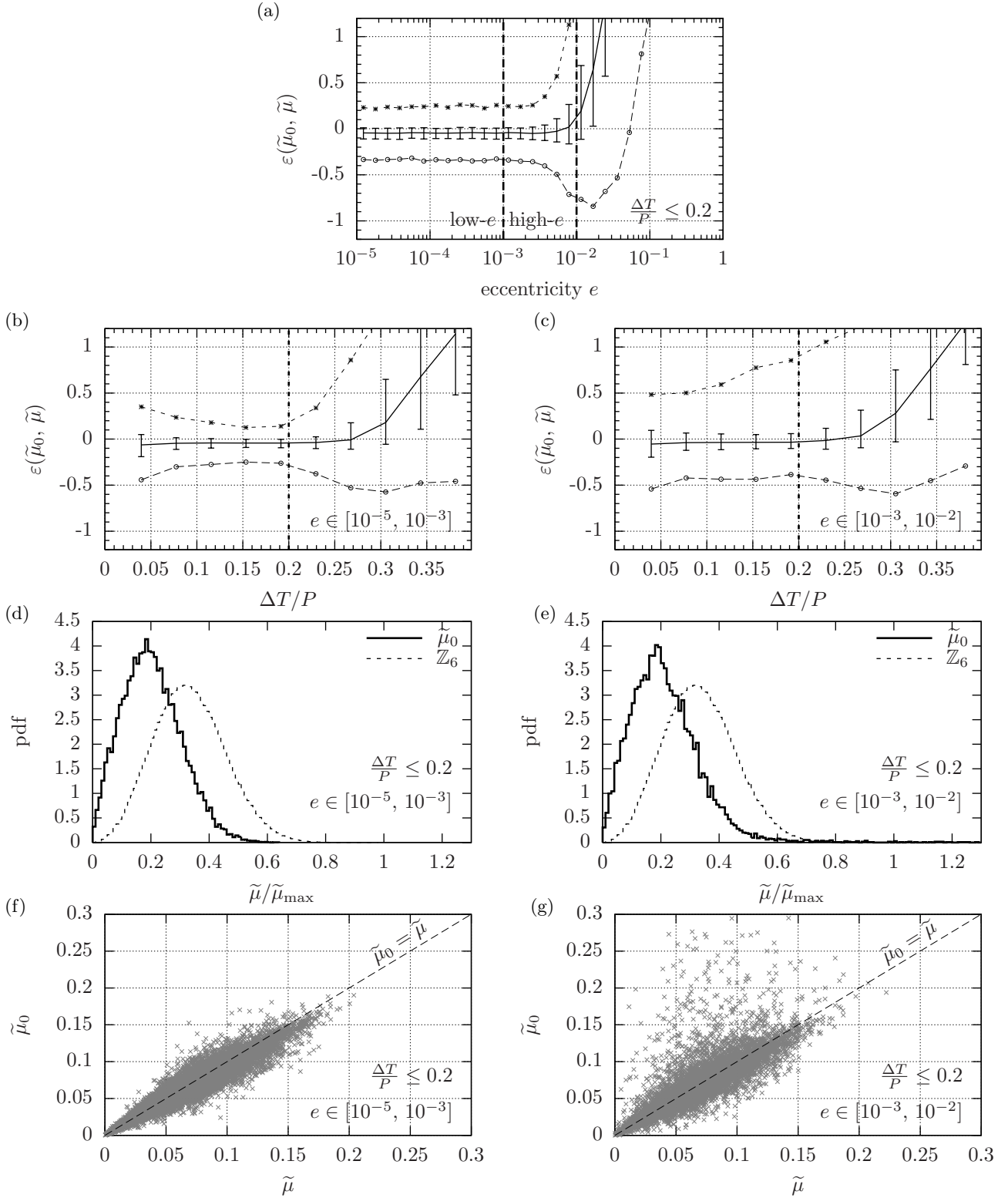


FIG. 4. *Coherent short-segment* regime: results for metric tests on \tilde{g}^{ss} of Sec. IV C 1 using u -coordinates, for orbital period scale of $P_0 = 80$ days and coherent observation times in the range $\Delta T \in [2, 25]$ days. (a) Relative error $\varepsilon(\tilde{\mu}_0, \tilde{\mu})$ versus eccentricity e . The dashed vertical line denotes the boundary between a “low- e ” range $e \in [10^{-5}, 10^{-3}]$ (left-column plots) and a “high- e ” range $e \in [10^{-3}, 10^{-2}]$ (right-column plots). (b,c): Relative error ε vs observation time $\Delta T/P$. The solid lines in panels (a, b, c) denote the median value, the error bars correspond to the 25th-75th percentiles, the circles and the stars denote the 2.5th and 97.5th percentiles, respectively. (d,e): Mismatch histogram of measured mismatches $\tilde{\mu}_0/\tilde{\mu}_{\text{max}}$ and theoretical distribution in a \mathbb{Z}_6 lattice. (f,g): Measured mismatch $\tilde{\mu}_0$ versus predicted $\tilde{\mu}$. Note that in all plots except (b,c) the range of relative segment length was restricted to $\Delta T/P \leq 0.2$, where the short-segment limit approximation is seen to be valid.

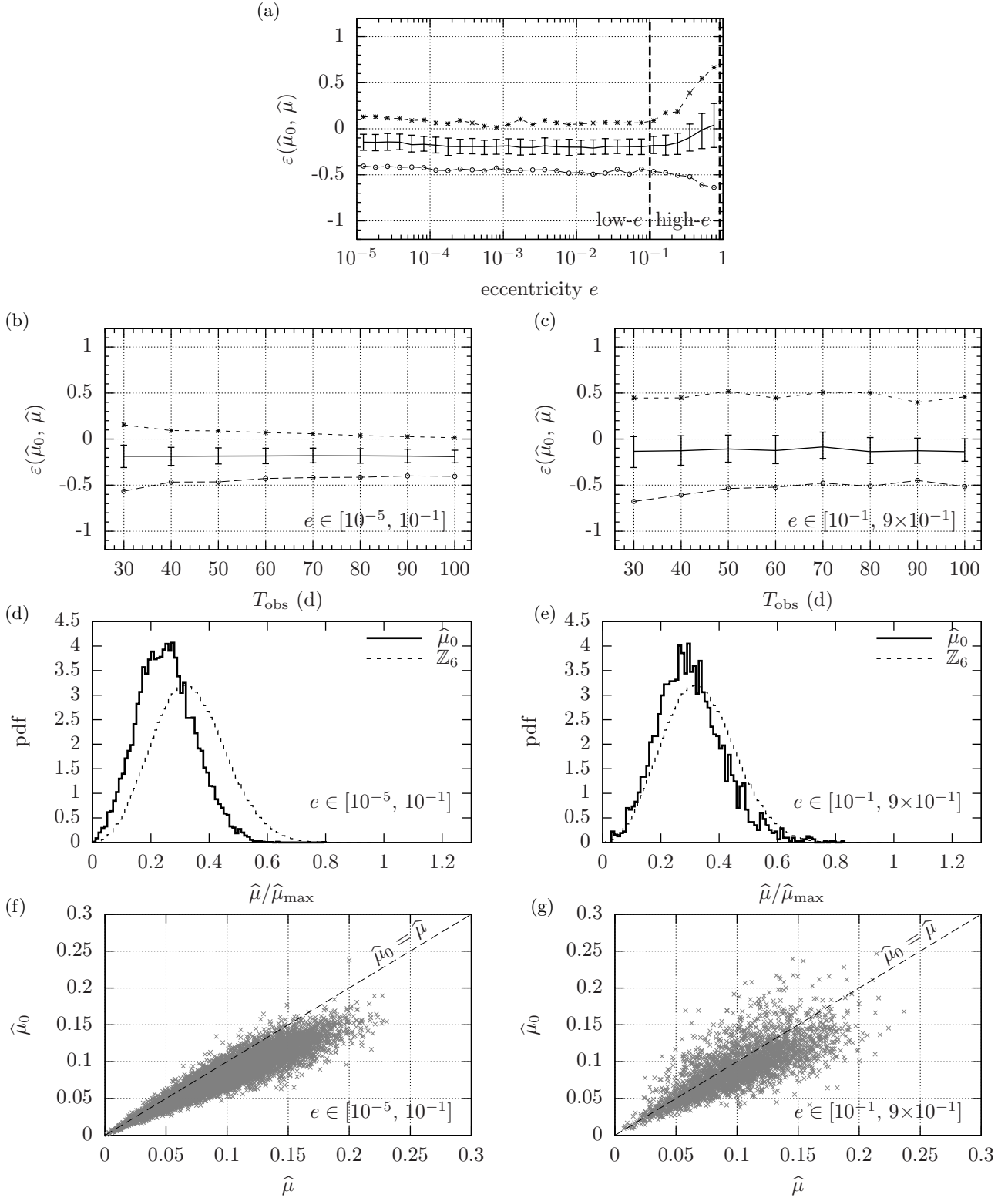


FIG. 5. *Semi-coherent short-segment* regime: results for metric tests on \hat{g}^{ss} of Sec. IVC2 for orbital period scale of $P_0 = 10$ days, and a fixed segment length of $\Delta T = 1$ day. (a) Relative error $\varepsilon(\hat{\mu}_0, \hat{\mu})$ versus eccentricity e . The dashed vertical line denotes the boundary between a “low- e ” range $e \in [10^{-5}, 0.1]$ (left-column plots) and a “high- e ” range $e \in [0.1, 0.9]$ (right-column plots). (b,c): Relative error ε vs observation time T_{obs} . The solid lines in panels (a,b,c) denote the median value, the error bars correspond to the 25th-75th percentiles, the circles and the stars denote the 2.5th and 97.5th percentiles, respectively. (d,e): Mismatch histogram of measured mismatches $\hat{\mu}_0/\hat{\mu}_{\text{max}}$ and theoretical distribution in a \mathbb{Z}_6 lattice. (f,g): Measured mismatch $\hat{\mu}_0$ versus predicted $\hat{\mu}$.

behaviour might be interesting, but is beyond the scope of this study.

The (phase-) metric approximation Eq. (6) is expected from previous studies [29, 30] to be quite accurate (in suitable coordinates) up to values of mismatch of $\mu \lesssim 0.2$, after which higher-order terms start to become more noticeable, which tend to *reduce* the actual measured mismatches compared to the metric predictions. While the nonlinear regime is hardly explored here, this general trend can still be seen to some extent in panels (f) of Figs. 1, 2, 4, 5.

VII. SCORPIUS X-1 SENSITIVITY ESTIMATE

We can use the metric expressions and template counts derived in Sec. V to estimate the optimal achievable sensitivity of a semi-coherent search directed at Scorpius X-1.

A. Torque-balance level

In order to quantify the sensitivity of a search *independent* of the detector noise floor, we can define the *sensitivity depth* [51] as

$$\mathcal{D}_{p_{\text{fa}}}^C \equiv \frac{\sqrt{S_n}}{h_{p_{\text{fa}}}^C}, \quad (84)$$

in terms of the harmonic mean over detector noise floors S_n , and the strain sensitivity (or upper limit) $h_{p_{\text{fa}}}^C$ of a search at a certain confidence level C (i.e. detection probability) and false-alarm threshold p_{fa} . This quantity (with dimensions of $\text{Hz}^{-1/2}$) is useful as a simple and intuitive measure for how far below the noise floor a given search setup can reach.

An interesting astrophysical model postulating torque balance between the accretion and CW emission [1, 2, 52] yields a predicted CW amplitude (assuming a NS with 10 km radius and a mass of $1.4 M_\odot$) for Scorpius X-1 of

$$h_0 \sim 3.5 \times 10^{-26} \sqrt{\frac{300 \text{ Hz}}{\nu}}, \quad (85)$$

where ν is the (unknown) NS spin frequency.

Figure 6 shows the minimum required sensitivity depth $\mathcal{D}^{\text{ScoX1}}$ to reach the Scorpius X-1 torque-balance limit of Eq. (85) for different emission models (“mountain” deformation with $\bar{f} = 2\nu$, r-mode emission with $\bar{f} \approx 4\nu/3$, and precession with $\bar{f} \approx \nu$), assuming different aLIGO sensitivity curves (“Early”, “Mid”, “Late” and “Final”) [53–55].

For all the following sensitivity estimates, we will assume a detector duty cycle of 80% (following [55]), a 90% confidence level (i.e. detection probability), and a false-alarm level of $p_{\text{fa}} = 10^{-10}$ (appropriate for a first-stage wide-parameter search).

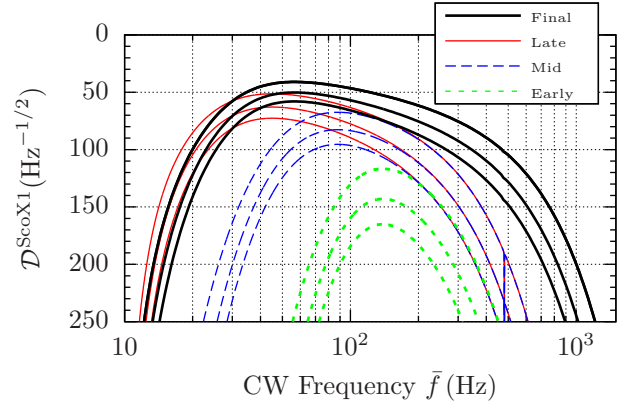


FIG. 6. Sensitivity depth $\mathcal{D}^{\text{ScoX1}}$ of Eq. (84) for Scorpius X-1 torque-balance level Eq. (85) versus CW frequency \bar{f} . Different line styles correspond to different progressive aLIGO sensitivities (“Early”, “Mid”, “Late” and “Final”) [53–55]. For each detector sensitivity three different emission models are shown, namely: “mountain” ($\bar{f} = 2\nu$), *mid*: r-mode ($\bar{f} \approx 4\nu/3$), *bottom*: precession ($\bar{f} \approx \nu$), respectively.

B. Best theoretically achievable sensitivity

Short-segment regime

Let us first consider the short-segment regime, assuming a “best-case” search with segments of length $\Delta T = 4.7 \text{ h}$ (i.e. roughly 1/4 of the Scorpius X-1 orbital period, quoted in Table I), total timespan $T_{\text{obs}} = 365 \text{ d}$ (therefore $N = 1864$ segments), 2 detectors [LIGO H and Livingston (L)], and a very fine search grid of average mismatch $\mu = 0.01$. Using the method outlined in [56] and implemented in [57], we obtain a resulting sensitivity depth of $\mathcal{D}_{1e-10}^{90\%} \sim 56 \text{ Hz}^{-1/2}$. Assuming 3 equal-sensitivity detectors [LIGO H, L and Virgo (V)], this value increases to $\mathcal{D}_{1e-10}^{90\%} \sim 68 \text{ Hz}^{-1/2}$.

Long-segment regime

In the long-segment regime of Sec. IV B, the maximal segment length would be restricted by the astrophysical concern of “spin wandering”, namely a stochastic variability of the spin frequency due to variations in the accretion rate. There is substantial astrophysical uncertainty [13, 18, 58] about the details of this process and its magnitude, which is beyond the scope of this work. We consider assumptions roughly similar to those given in the recent Scorpius X-1 MDC [10], which posited a random frequency derivative of order $|\dot{f}| \leq \dot{f}_s \sim 10^{-12} \text{ Hz/s}$, changing on a timescale of $t_s \sim 10^6 \text{ s}$. Over this timescale, the frequency drift would therefore be $\delta \tilde{f}_s \leq \dot{f}_s t_s \sim 10^{-6} \text{ Hz}$. Over a total observation time T_{obs} , this can be modelled as a random walk with expected total drift (with respect to the mid-

TABLE I. Assumed Scorpius X-1 system parameters and uncertainties. Ranges of the form $\lambda \pm \Delta\lambda$ denote the mean and Gaussian 1-sigma uncertainty, while $[\lambda_0, \lambda_1]$ denotes a uniform probability range. The time of ascending node t_{asc} was computed from the originally-measured time of inferior conjunction T_0 [48] via $t_{\text{asc}} = T_0 - P/4$. The projected semi-major axis a_p is related to the projected radial maximal velocity K_1 , appearing in [60], by $a_p = K_1 P/(2\pi c)$.

Parameter	Value	Ref.
a_p (s)	1.44 ± 0.18	[13, 60]
t_{asc} (GPS s)	897753994 ± 100	[48]
P (s)	68023.70496 ± 0.0432	[48]
e	(≈ 0) or (0.033 ± 0.018)	[59]
ω (rad)	$[0, 2\pi]$	[59]

point in time) of order $\delta\hat{f}_s \sim \sqrt{T_{\text{obs}}/(2t_s)} \delta\tilde{f}_s$. The maximal coherent segment length could therefore be estimated roughly by the requirement that the total frequency drift should be less than the frequency resolution $\sim 1/\Delta T$ of the search, in order to avoid any significant loss of SNR. This yields the constraint

$$\Delta T \lesssim \sqrt{2} \left(\sqrt{T_{\text{obs}} t_s} \dot{f}_s \right)^{-1}. \quad (86)$$

Assuming $T_{\text{obs}} = 1$ year and the above MDC spin-wandering model, we find $\Delta T \lesssim 3$ d. Given the substantial uncertainty in these parameters, for comparison we also consider a more optimistic scenario of $\Delta T \lesssim 10$ d.

The best achievable sensitivity for $\Delta T = 3$ d and $N = 120$ segments, assuming a small mismatch of $\mu = 0.01$, can be estimated for a 2-detector network as $\mathcal{D}_{1e-10}^{90\%} \sim 105 \text{ Hz}^{-1/2}$, and for 3 detectors we find $\mathcal{D}_{1e-10}^{90\%} \sim 127 \text{ Hz}^{-1/2}$.

The best sensitivity that can be achieved for $\Delta T = 10$ d and $N = 36$ segments, and $\mu = 0.01$, can be estimated for a 2-detector network as $\mathcal{D}_{1e-10}^{90\%} \sim 134 \text{ Hz}^{-1/2}$, and for 3 detectors we find $\mathcal{D}_{1e-10}^{90\%} \sim 163 \text{ Hz}^{-1/2}$.

In the following we will focus only on the long-segment regime, and assess the required computing power as well as the resulting sensitivity depth that can actually be reached in practice.

C. Scorpius X-1 search parameter space

We assume the Scorpius X-1 parameters and uncertainties given in Table I. Note that, contrary to the Scorpius X-1 MDC [10] and previous searches [13], we also allow for a nonzero uncertainty on the eccentricity, in addition to the circular-orbit assumption that is also not ruled out by current observations [59].

Assuming 3-sigma ranges for the Scorpius X-1 parameter uncertainties, the resulting per-dimension template numbers \mathcal{N}_{λ_i} at maximal mismatch $\mu_{\text{max}} = 0.1$ of Eq. (70) are shown in Fig. 7. We see that, for the poten-

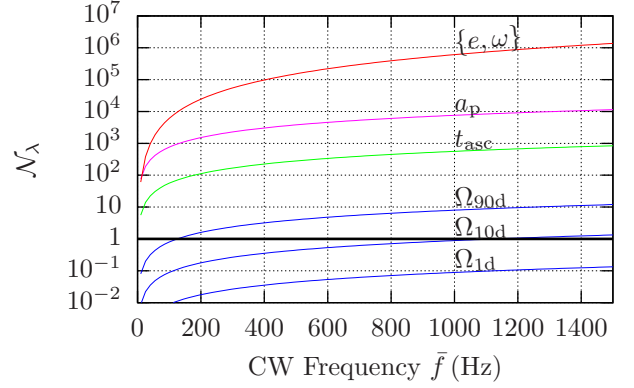


FIG. 7. Effective number of (coherent or semi-coherent) templates per dimension \mathcal{N}_{λ} of Eq. (70) (in the long-segment regime) as function of CW frequency \bar{f} for the Scorpius X-1 parameter space given in Table I, using ± 3 sigma ranges. The template numbers are evaluated for $\mu_{\text{max}} = 0.1$ and mean values of a_p and Ω . Note that assuming a 2π uncertainty in ω entails $\mathcal{N}_{\{e, \omega\}} = \mathcal{N}_e^2 \pi$. The template count \mathcal{N}_{Ω} is the only one depending on the observation time T and is shown for $T = 1$ d, 10 d, 90 d.

tial uncertainty on eccentricity e of Table I, one would have to include e and ω as effective search dimensions in the long-segment limit. On the other hand, the assumption of full circularization $e \approx 0$ is not ruled out by the current observations and therefore also constitutes a reasonable alternative model to investigate.

The only dimension where the resolution depends on the timespan T is Ω . We see, for example, that for coherent times $\Delta T \lesssim 10$ d the uncertainty in Ω will typically not be resolved by the coherent metric (as long as $\tilde{\mu}_{\text{max}} \geq 0.1$), as $\mathcal{N}_{\Omega}(T < 10\text{d}) \lesssim 1$. In this case one could simply exclude this search dimension in the coherent template count $\tilde{\mathcal{N}}$ of Eq. (13). On the other hand, the semi-coherent fine grid (assuming $T_{\text{obs}} \geq 90$ d and $\hat{\mu}_{\text{max}} \leq 0.1$) can typically resolve this dimension over some frequency range where $\mathcal{N}_{\Omega}(T > 90\text{d}) > 1$. In both cases, however, this depends on the resulting mismatch parameter μ_{max} , which is itself an optimization parameter. Therefore, as seen in Fig. 7, one potentially has to deal with situations where the Ω -dimension will be unresolved at lower frequencies but resolved at higher frequencies. One way to obtain a correct template estimate in such a case would be to determine if a transition frequency exists, where the Ω -dimension changes from unresolved to resolved, then stitch together the respective correct template counts over each frequency range. Alternatively, one can simply split the frequency range into narrow slices, such that the dimensionality can be assumed constant over each slice, and sum up the respective number of templates.

Note that due to the large separation Δ_{ma} between the orbital epoch t_{asc} in Table I and the aLIGO gravitational wave search epoch, we have to consider the potential increase in uncertainty on t_{asc} , as discussed in [48]. As dis-

cussed in Sec. V A, this is not a concern for a grid where both Ω and t_{asc} are fully resolved by the metric, as is expected to be the case of the semi-coherent grid. On the other hand, the coherent grids will typically not resolve the Ω dimension, and therefore an increase of about 2.5–3 in uncertainty on t_{asc} would have to be taken into account. However, for the present study, we assume (as seems likely [48]) that there will be further observations on Scorpius X-1 closer to the aLIGO epoch, which will again constrain the uncertainty on t_{asc} to a level similar (or better) than that given in Table I.

D. Computing cost model

As described in [20], a semi-coherent StackSlide search (introduced in Sec. II B) for a fixed parameter space has the following tuneable parameters: the segment length ΔT , the number of segments N , the maximal template-bank mismatch for the (semi-coherent) fine grid $\hat{\mu}_{\text{max}}$, and for the per-segment (coherent) coarse grids $\tilde{\mu}_{\text{max}}$. The aim is therefore to maximize the resulting sensitivity over this space of search parameters, under the constraint of fixed total computing cost $\mathcal{C}_{\text{tot}}(\tilde{\mu}_{\text{max}}, \hat{\mu}_{\text{max}}, N, \Delta T) = \mathcal{C}_0$.

The total computing cost \mathcal{C}_{tot} of the StackSlide $\hat{\mathcal{F}}$ -statistic can be written as [20]

$$\mathcal{C}_{\text{tot}}(\tilde{\mu}_{\text{max}}, \hat{\mu}_{\text{max}}, N, \Delta T) = \tilde{\mathcal{C}} + \hat{\mathcal{C}}, \quad (87)$$

where $\tilde{\mathcal{C}}(\tilde{\mu}_{\text{max}}, N, \Delta T)$ is the cost of N per-segment coherent $\tilde{\mathcal{F}}$ -statistic searches over a coarse grid with maximal mismatch $\tilde{\mu}_{\text{max}}$, and $\hat{\mathcal{C}}(\hat{\mu}_{\text{max}}, N, \Delta T)$ is the cost of incoherently summing (and interpolating) these $\tilde{\mathcal{F}}_\ell$ -values across all segments on a fine grid with maximal mismatch $\hat{\mu}_{\text{max}}$.

The coherent cost for the $\tilde{\mathcal{F}}$ -statistic computation can be expressed as

$$\tilde{\mathcal{C}} = N \tilde{\mathcal{N}}(\tilde{\mu}_{\text{max}}, \Delta T) N_{\text{det}} \tilde{c}_1(\Delta T), \quad (88)$$

where N_{det} is the number of detectors, and \tilde{c}_1 is the (per-detector) computing cost per template, which depends on the algorithm used to compute the $\tilde{\mathcal{F}}$ -statistic. For the (generally slower) SFT-based demodulation method [61], this can be expressed (for gapless data) as

$$\tilde{c}_1^{\text{sft}}(\Delta T) = \tilde{c}_0^{\text{sft}} \frac{\Delta T}{T_{\text{sft}}}, \quad (89)$$

where $\tilde{c}_0^{\text{sft}} \approx 4 \times 10^{-8}$ s is an implementation- and hardware-dependent fundamental computing cost for binary-CW templates (measured by timing the current $\tilde{\mathcal{F}}$ -statistic LALSuite code [49]).

Using instead the more efficient Fast Fourier Transform (FFT) based “resampling” method [28, 62], we find the time per template \tilde{c}_1 to be approximately constant (for search frequency bands of $\gtrsim 10^5$ frequency bins) as

$$\tilde{c}_1^{\text{FFT}} \approx 3 \times 10^{-7} \text{ s}. \quad (90)$$

The semi-coherent cost of summing across segments can be expressed as

$$\hat{\mathcal{C}} = N \hat{\mathcal{N}}(\hat{\mu}_{\text{max}}, \Delta T, N) \hat{c}_0, \quad (91)$$

where $\hat{c}_0 \approx 5 \times 10^{-9}$ s is an implementation- and hardware-dependent fundamental computing cost of adding one value of $2 \tilde{\mathcal{F}}_\ell$ for one fine-grid point (measured by timing the LALSuite `HierarchSearchGCT` code [49]).

Note that the sensitivity estimates presented in Tables II–VII are using “Einstein@Home months” (EM) as a computing-cost unit. This unit is defined as 12 000 core-months⁷ on a CPU that achieves the above-quoted timings \hat{c}_0 and \tilde{c}_1 . Empirically this corresponds to the performance of the “average CPU” currently participating in Einstein@Home, and is comparable to roughly an Intel Core i7-2620M or Intel Xeon E3-1220 v3.

E. Sensitivity estimates

With the computing-cost model at hand we can employ the semi-analytic machinery of [20] to find optimal StackSlide setups for given computing cost. We can then estimate the corresponding sensitivity of each setup with the fast and accurate method developed in [56]. Both the optimization and sensitivity-estimation algorithms have been implemented in [57] and are used here.

Due to the rapidly increasing computing cost with frequency [see Eqs. (72)–(75)], it is important to limit the frequency range searched in some way. Hence, we focus on constructing example setups that beat the torque-balance level (shown in Fig. 6) over their respective frequency range.

We consider the two different constraints on the segment length ΔT due to spin wandering, as discussed in Sec. VII B, namely $\Delta T \leq 10$ d and $\Delta T \leq 3$ d, respectively. For $\Delta T \leq 10$ d, the results for different \mathcal{F} -computation methods and assumptions about eccentricity are summarized in Table II (for SFT-based demodulation and zero uncertainty on e), Table III (for FFT-based resampling and zero uncertainty on e), and Table IV (for FFT-based resampling and 3-sigma uncertainty of $\Delta e = 0.087$).

Similarly, for $\Delta T \leq 3$ d, the results for different \mathcal{F} -computation methods and assumptions about eccentricity are summarized in Table V (for SFT-based demodulation and zero uncertainty on e), Table VI (for FFT-based resampling and zero uncertainty on e), and Table VII (for FFT-based resampling and 3-sigma uncertainty of $\Delta e = 0.087$).

In each table the first line shows a setup beating the $\bar{f} = 2\nu$ torque-balance depth for the “Final” aLIGO sensitivity curve of Fig. 6 over the given frequency range,

⁷ Counting only the 30%-50% CPUs currently devoted to gravitational wave searches, and including the double-computation of all results for validation.

while subsequent lines illustrate slight variations when changing only one of the constraints (computing cost \mathcal{C}_0 , observation time T_{obs} or detectors, respectively). In the last line of each table we also show a setup beating the $\bar{f} = 2\nu$ torque-balance level depth over the given frequency range assuming a 6-month science run using the “Mid” aLIGO configuration, which is currently estimated [55] to take place in about 2016-2017.

Summarizing these results, we find that the SFT-based demodulation method can only reach the torque-balance level assuming we can neglect the uncertainty in eccentricity ($\Delta e = 0$), shown in Tables II and V. The FFT-based resampling method, in the case of $\Delta e = 0$, can beat the torque-balance level over a wider frequency range up to ~ 630 Hz (for $\Delta T \leq 10$ d) or up to ~ 500 Hz (for $\Delta T \leq 3$ d), even at substantially reduced computing cost (cf. Tables III and VI), and can still beat the torque-balance limit (albeit only up to a smaller frequency of ~ 160 Hz – 200 Hz, depending on the constraint on ΔT) in the case of substantial uncertainty ($\Delta e = 0.087$) on eccentricity, as seen in Tables IV and VII.

At a lower-false alarm probability of $p_{\text{fa}} = 10^{-14}$, corresponding to a level at which one might set upper limits without any further followup, the corresponding sensitivity depth $\mathcal{D}_{1e-14}^{90\%}$ would be roughly 10% lower compared to the values of $\mathcal{D}_{1e-10}^{90\%}$ quoted in Tables II– VII.

Note that some of the optimal setups include maximal mismatches μ_{max} of order unity, which would violate the metric approximation discussed in Sec. II C. In order to fully quantify this effect, further Monte-Carlo tests on the higher-order metric deviations would be required, but qualitatively we know from previous results [e.g. see [29, 30], Figs. 1 (f), 2 (f), and metric tests we performed at $\mu_{\text{max}} = 0.5$ that are not shown here] that in this regime the metric approximation generally tends to *overestimate* the actual measured mismatches. Furthermore, the distribution of sampled mismatches will be peaked around their mean value, which for A_4^* lattices is $\langle \mu \rangle \approx 0.5 \mu_{\text{max}}$, and will therefore not be affected as much by these deviations.

F. Comparison to previous sensitivities

In order to put these sensitivities in context, we consider some previously-achieved sensitivities for Scorpius X-1 CW searches. The first coherent Scorpius X-1 search presented in [11] was computationally limited to using a total of ~ 9 h of data collected by LIGO during the second science run with the Livingston and Hanford detectors. Such a search achieved upper limits corresponding to a sensitivity depth of approximately $\mathcal{D}^{95\%} \approx 4 \text{ Hz}^{-1/2}$, which is roughly consistent with the theoretical expectation [56] of $\mathcal{D}_{1e-14}^{95\%} \approx 4.5 \text{ Hz}^{-1/2}$ for a fully-coherent two-detector search. The recent search on the fifth LIGO science run, which uses a semi-coherent sideband method [13], achieved a substantially improved

sensitivity depth of $\mathcal{D}^{95\%} \approx 30 \text{ Hz}^{-1/2}$ over a frequency band of 50 – 550 Hz. Note that the software-injected signals used in the recent Scorpius X-1 MDC [10] are found (over a frequency range of 50 – 1500 Hz) in the depth range $\mathcal{D} \in [2.6, 104] \text{ Hz}^{-1/2}$ for the “open” injections, and $\mathcal{D} \in [2.4, 58] \text{ Hz}^{-1/2}$ for the “closed” injections.

VIII. CONCLUSIONS

The work presented here consists of three main parts.

In the first part we re-derived the (coherent and semi-coherent) metric expressions for the long-segment and short-segment regimes, assuming a low-eccentricity binary orbit. We found that (as has also been noted earlier), the long-segment regime does require refinement in the orbital angular velocity Ω , contrary to an earlier result in the literature [27]. We have extended these findings to allow for general offsets between the orbital reference time epoch t_{asc} and the observation epoch, which explicitly show that there is generally *no* increase in computing cost with increasing time passed since the original orbital reference epoch if both Ω and t_{asc} are fully resolved by the metric.

In the second part we subjected the analytic metric expressions to extensive Monte-Carlo tests comparing their predictions against measured SNR loss in software-injection studies, and found robust agreement within their range of applicability.

In the final part of this work we used the metric template counts to semi-analytically estimate the optimal achievable StackSlide sensitivity of a directed search for Scorpius X-1. We found that the predicted torque-balance limit could be reachable for the first time with an Einstein@Home search using data from a 6-month “Mid” aLIGO run (currently planned for 2016-2017), and should be beatable over a substantial frequency range in “Final” aLIGO. However, the frequency range over which the torque-balance limit can be beaten depends strongly on the Scorpius X-1 parameter-space uncertainty, most notably the orbital eccentricity and the assumed characteristics of spin wandering.

More work is required to better understand the modelling and effects of spin wandering, as well as the resulting constraints on search methods. It would seem desirable to develop a more robust statistic that properly takes spin wandering into account, for example by marginalizing over the spindown uncertainty in each segment before incoherently summing them.

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TABLE II. Optimal search setups and corresponding sensitivity estimates for a Scorpius X-1 StackSlide search (with $\Delta T \leq 10$ d) using the SFT-based demodulation algorithm, assuming negligible uncertainty $\Delta e = 0$ on the eccentricity. The SFT length assumed is $T_{\text{sft}}(430 \text{ Hz}) = 240 \text{ s}$ and $T_{\text{sft}}(250 \text{ Hz}) = 320 \text{ s}$ (see Appendix C, for $\mu_{\text{sft}} = 0.01$). The first line shows a setup beating the “Final” $\bar{f} = 2\nu$ torque-balance depth (see Fig. 6) over the given frequency range, and subsequent lines illustrate the results for slightly different constraints on computing cost \mathcal{C}_0 , observation time T_{obs} and detectors (IFOs). The last line shows a setup beating the $\bar{f} = 2\nu$ torque-balance level depth over the given frequency range with a “Mid” aLIGO configuration. In all cases the template-bank dimensionalities are found as $\tilde{n} = 3$ and $\hat{n} = 4$. The computing-cost unit “EM” corresponds to running on Einstein@Home for a month, and is explained in more detail in Sec. VII D.

\bar{f} (Hz)	\mathcal{C}_0 (EM)	IFOs	T_{obs} (d)	N	ΔT (d)	$\tilde{\mu}_{\text{max}}$	$\hat{\mu}_{\text{max}}$	$\tilde{\mathcal{N}}$	$\hat{\mathcal{N}}$	$\tilde{\mathcal{C}}/\hat{\mathcal{C}}$	$\mathcal{D}_{1\text{e}-10}^{90\%}$ ($1/\sqrt{\text{Hz}}$)
[20, 430]	12	HL	360.0	43	8.30	0.71	0.04	3.4×10^{13}	7.5×10^{16}	22	100
[20, 430]	6	HL	360.0	83	4.36	0.74	0.06	1.7×10^{13}	2.4×10^{16}	18	86
[20, 430]	12	HL	180.0	18	10.00	0.50	0.02	6.9×10^{13}	1.5×10^{17}	27	91
[20, 430]	12	HLV	360.0	62	5.77	0.73	0.04	2.3×10^{13}	5.1×10^{16}	22	112
6 months science run with “Mid” aLIGO:											
[40, 230]	6	HL	180.0	18	10.00	0.18	0.010	4.8×10^{13}	7.8×10^{16}	26	102

TABLE III. Optimal search setups and corresponding sensitivity estimates for a Scorpius X-1 StackSlide search (with $\Delta T \leq 10$ d) using the FFT-based resampling algorithm, assuming negligible uncertainty $\Delta e = 0$ on the eccentricity. In all cases the template-bank dimensionalities are found as $\tilde{n} = 3$ and $\hat{n} = 4$. The notation is the same as in Table II.

\bar{f} (Hz)	\mathcal{C}_0 (EM)	IFOs	T_{obs} (d)	N	ΔT (d)	$\tilde{\mu}_{\text{max}}$	$\hat{\mu}_{\text{max}}$	$\tilde{\mathcal{N}}$	$\hat{\mathcal{N}}$	$\tilde{\mathcal{C}}/\hat{\mathcal{C}}$	$\mathcal{D}_{1\text{e}-10}^{90\%}$ ($1/\sqrt{\text{Hz}}$)
[20, 630]	12	HL	360.0	36	10.00	0.04	0.03	1.1×10^{16}	8×10^{17}	1.6	129
[20, 630]	6	HL	360.0	36	10.00	0.06	0.05	5.5×10^{15}	3.8×10^{17}	1.7	127
[20, 630]	1	HL	360.0	36	10.00	0.19	0.12	9.7×10^{14}	5.6×10^{16}	2.1	117
[20, 630]	12	HL	180.0	18	10.00	0.03	0.02	2.5×10^{16}	1.4×10^{18}	2.1	105
[20, 630]	12	HLV	360.0	36	10.00	0.05	0.03	7.5×10^{15}	7.2×10^{17}	1.9	155
6 months science run with “Mid” aLIGO:											
[40, 250]	6	HL	180.0	18	10.00	0.006	0.004	1.1×10^{16}	6.9×10^{17}	2	107

TABLE IV. Optimal search setups and corresponding sensitivity estimates for a Scorpius X-1 StackSlide search (with $\Delta T \leq 10$ d) using the FFT-based resampling method, assuming a pessimistic uncertainty $\Delta e = 0.087$ on eccentricity (see Table I). In all cases the template-bank dimensionalities are found as $\tilde{n} = 5$ and $\hat{n} = 6$. The notation is the same as in Table II.

\bar{f} (Hz)	\mathcal{C}_0 (EM)	IFOs	T_{obs} (d)	N	ΔT (d)	$\tilde{\mu}_{\text{max}}$	$\hat{\mu}_{\text{max}}$	$\tilde{\mathcal{N}}$	$\hat{\mathcal{N}}$	$\tilde{\mathcal{C}}/\hat{\mathcal{C}}$	$\mathcal{D}_{1\text{e}-10}^{90\%}$ ($1/\sqrt{\text{Hz}}$)
[20, 200]	12	HL	360.0	36	10.00	1.18	0.54	1.3×10^{16}	5.7×10^{17}	2.6	60
[20, 200]	6	HL	360.0	36	10.00	1.56	0.68	6.3×10^{15}	2.8×10^{17}	2.7	55
[20, 200]	12	HL	180.0	18	10.00	0.88	0.35	2.6×10^{16}	10^{18}	3	56
[20, 200]	12	HLV	360.0	36	10.00	1.37	0.55	8.6×10^{15}	5.2×10^{17}	3	68
6 months science run with “Mid” aLIGO:											
[60, 140]	6	HL	180.0	18	10.00	0.57	0.21	1.3×10^{16}	5×10^{17}	3.1	74

unpublished parameter estimate on the eccentricity of Scorpius X-1. PL thanks the support of the *Sonderforschungsbereich Transregio* (SFB/TR7) collaboration and the University of Rome “Sapienza”. Numerical simulations were performed on the ATLAS computer cluster of the *Max-Planck-Institut für Gravitationsphysik*. This work was completed using exclusively FreeSoftware tools: the algebra was checked using Maxima, the numerical scripting was done in Octave with SWIG bindings (pro-

vided by Karl Wette) to LALSuite, and all plots were prepared using Octave/Gnuplot. This paper has been assigned document number LIGO-P1500006-v3.

TABLE V. Optimal search setups and corresponding sensitivity estimates for a Scorpius X-1 StackSlide search (with $\Delta T \leq 3$ d) using the SFT-based demodulation algorithm, assuming negligible uncertainty $\Delta e = 0$ on the eccentricity. The SFT length assumed is $T_{\text{sft}}(420\text{Hz}) = 250$ s and $T_{\text{sft}}(150\text{Hz}) = 410$ s (see Appendix C, for $\mu_{\text{sft}} = 0.01$). The first line shows a setup beating the “Final” $\bar{f} = 2\nu$ torque-balance depth (see Fig. 6) over the given frequency range, and subsequent lines illustrate the results for slightly different constraints on computing cost \mathcal{C}_0 , observation time T_{obs} and detectors (IFOs). The last line shows a setup beating the $\bar{f} = 2\nu$ torque-balance level depth over the given frequency range with a “Mid” aLIGO configuration. In all cases the template-bank dimensions are found as $\tilde{n} = 3$ and $\hat{n} = 4$.

\bar{f} (Hz)	\mathcal{C}_0 (EM)	IFOs	T_{obs} (d)	N	ΔT (d)	$\tilde{\mu}_{\text{max}}$	$\hat{\mu}_{\text{max}}$	$\tilde{\mathcal{N}}$	$\hat{\mathcal{N}}$	$\tilde{\mathcal{C}}/\hat{\mathcal{C}}$	$\mathcal{D}_{1e-10}^{90\%}$ ($1/\sqrt{\text{Hz}}$)
[20, 420]	12	HL	360.0	120	3.00	0.34	0.03	3.5×10^{13}	4.3×10^{16}	14	94
[20, 420]	6	HL	360.0	120	3.00	0.54	0.05	1.8×10^{13}	2×10^{16}	15	87
[20, 420]	12	HL	180.0	60	3.00	0.21	0.02	7.1×10^{13}	7.4×10^{16}	16	80
[20, 420]	12	HLV	360.0	120	3.00	0.44	0.04	2.4×10^{13}	3.6×10^{16}	16	110
6 months science run with “Mid” aLIGO:											
[50, 150]	6	HL	180.0	60	3.00	0.03	0.003	5.7×10^{13}	4.3×10^{16}	14	85

TABLE VI. Optimal search setups and corresponding sensitivity estimates for a Scorpius X-1 StackSlide search (with $\Delta T \leq 3$ d) using the FFT-based resampling algorithm, assuming negligible uncertainty $\Delta e = 0$ on the eccentricity. In all cases the template-bank dimensionalities are found as $\tilde{n} = 3$ and $\hat{n} = 4$. The notation is the same as in Table II.

\bar{f} (Hz)	\mathcal{C}_0 (EM)	IFOs	T_{obs} (d)	N	ΔT (d)	$\tilde{\mu}_{\text{max}}$	$\hat{\mu}_{\text{max}}$	$\tilde{\mathcal{N}}$	$\hat{\mathcal{N}}$	$\tilde{\mathcal{C}}/\hat{\mathcal{C}}$	$\mathcal{D}_{1e-10}^{90\%}$ ($1/\sqrt{\text{Hz}}$)
[20, 500]	12	HL	360.0	120	3.00	0.02	0.02	3.2×10^{15}	2.4×10^{17}	1.6	102
[20, 500]	6	HL	360.0	120	3.00	0.04	0.03	1.6×10^{15}	1.1×10^{17}	1.7	101
[20, 500]	1	HL	360.0	120	3.00	0.12	0.07	2.9×10^{14}	1.7×10^{16}	2.1	96
[20, 500]	12	HL	180.0	60	3.00	0.01	0.01	6.7×10^{15}	4.4×10^{17}	1.8	84
[20, 500]	12	HLV	360.0	120	3.00	0.03	0.02	2.3×10^{15}	2.1×10^{17}	1.9	123
6 months science run with “Mid” aLIGO:											
[50, 190]	6	HL	180.0	60	3.00	0.003	0.002	3.4×10^{15}	2.1×10^{17}	2	85

TABLE VII. Optimal search setups and corresponding sensitivity estimates for a Scorpius X-1 StackSlide search (with $\Delta T \leq 3$ d) using the FFT-based resampling method, assuming a pessimistic uncertainty $\Delta e = 0.087$ on eccentricity (see Table I). In all cases the template-bank dimensionalities are found as $\tilde{n} = 5$ and $\hat{n} = 6$. The notation is the same as in Table II.

\bar{f} (Hz)	\mathcal{C}_0 (EM)	IFOs	T_{obs} (d)	N	ΔT (d)	$\tilde{\mu}_{\text{max}}$	$\hat{\mu}_{\text{max}}$	$\tilde{\mathcal{N}}$	$\hat{\mathcal{N}}$	$\tilde{\mathcal{C}}/\hat{\mathcal{C}}$	$\mathcal{D}_{1e-10}^{90\%}$ ($1/\sqrt{\text{Hz}}$)
[20, 160]	12	HL	360.0	120	3.00	0.76	0.34	3.8×10^{15}	1.7×10^{17}	2.6	58
[20, 160]	6	HL	360.0	120	3.00	1.00	0.44	1.9×10^{15}	8.3×10^{16}	2.7	50
[20, 160]	12	HL	180.0	60	3.00	0.57	0.22	7.8×10^{15}	3.1×10^{17}	3	59
[20, 160]	12	HLV	360.0	120	3.00	0.88	0.35	2.6×10^{15}	1.6×10^{17}	3	65
6 months science run with “Mid” aLIGO:											
[70, 110]	6	HL	180.0	60	3.00	0.34	0.13	3.9×10^{15}	1.5×10^{17}	3.1	71

Appendix A: Inverting the u -coordinates back into physical coordinates

1. General 6D elliptic case

Let us define re-scaled coordinates

$$w_k \equiv \frac{u_k}{\Omega^k}, \quad (\text{A1})$$

and consider the set of Eqs. (58) for $w_k = w_k(\lambda)$.

Let us take linear combinations of the equations for w_2, \dots, w_5 to express

$$\sin \psi_m = \frac{1}{3\bar{f}a_p} (w_4 + 4w_2), \quad (\text{A2})$$

$$\cos \psi_m = \frac{1}{3\bar{f}a_p} (w_5 + 4w_3), \quad (\text{A3})$$

$$\sin 2\psi_m = -\frac{1}{12\bar{f}a_p e^2} [2\kappa(w_2 + w_4) + \eta(w_3 + w_5)], \quad (\text{A4})$$

$$\cos \psi_m = -\frac{1}{12\bar{f}a_p e^2} [\kappa(w_3 + w_5) - 2\eta(w_2 + w_4)], \quad (\text{A5})$$

where we used $e^2 = \kappa^2 + \eta^2$. We now insert these into the equation for w_6 to obtain

$$w_6 + 5w_4 + 4w_2 = 0, \quad (\text{A6})$$

which yields a quadratic equation for Ω^2 , with solution

$$\Omega^2 = -\frac{5u_4}{8u_2} + \sqrt{\left(\frac{5u_4}{8u_2}\right)^2 - \frac{u_6}{4u_2}}, \quad (\text{A7})$$

which requires $25u_4^2 - 16u_2u_6 > 0$ for real-valued Ω^2 . Note that once we know Ω , then all w_k are known numerically as well. Inserting Eqs. (A2)-(A5) into the equation for w_1 brings us to find

$$f = u_1 + \frac{5u_3}{\Omega^2} + \frac{u_5}{4\Omega^4}. \quad (\text{A8})$$

Using $\sin^2 \psi_m + \cos^2 \psi_m = 1$, we obtain

$$3\bar{f}a_p = \sqrt{(w_5 + 4w_3)^2 + (w_4 + 4w_2)^2}, \quad (\text{A9})$$

where we note that for constant-frequency signals, we numerically have $\bar{f} = f$. Hence, given f from the previous equation, we easily get a_p . Proceeding similarly for $\sin^2 2\psi_m + \cos^2 2\psi_m = 1$, we find

$$e = \frac{1}{12\bar{f}a_p} \sqrt{(w_3 + w_5)^2 + 4(w_2 + w_4)^2}. \quad (\text{A10})$$

We can simply invert Eqs. (A2),(A3) to obtain ψ_m and therefore t_{asc} . Finally, solving the linear system of equations w_2, w_3 for κ, η , yields

$$\kappa = \frac{1}{4\bar{f}a_p} [2(w_2 - \bar{f}a_p \sin \psi_m) \sin 2\psi_m + (w_3 - \bar{f}a_p \cos \psi_m) \cos 2\psi_m], \quad (\text{A11})$$

$$\eta = \frac{1}{4\bar{f}a_p} [(w_3 - \bar{f}a_p \cos \psi_m) \sin 2\psi_m + 2(\bar{f}a_p \sin \psi_m - w_2) \cos 2\psi_m], \quad (\text{A12})$$

which, using Eq. (A10) and Eq. (33) gives us ω .

2. Special 4D circular case

In the special circular case, we can obtain a simpler solution by inverting the equations for $\{w_1, \dots, w_4\}$ and setting $\kappa = \eta = 0$, namely

$$w_1 = \frac{f}{\Omega} - \bar{f}a_p \cos \psi_m, \quad (\text{A13})$$

$$w_2 = \bar{f}a_p \sin \psi_m, \quad (\text{A14})$$

$$w_3 = \bar{f}a_p \cos \psi_m, \quad (\text{A15})$$

$$w_4 = -\bar{f}a_p \sin \psi_m. \quad (\text{A16})$$

By summing equations for w_2 and w_4 , we find $w_2 + w_4 = 0$, which is a quadratic equation for Ω with solution

$$\Omega = \sqrt{-\frac{u_4}{u_2}}, \quad (\text{A17})$$

which numerically determines all w_k . We can further see that

$$\bar{f}a_p = \sqrt{w_2^2 + w_3^2}, \quad (\text{A18})$$

and

$$\tan \psi_m = \frac{w_2}{w_3}, \quad (\text{A19})$$

yielding t_{asc} , and finally

$$f = \bar{f}a_p \Omega \cos \psi_m + u_1. \quad (\text{A20})$$

Appendix B: Maximal Doppler shift due to orbital motion

Sometimes it is important to estimate the maximal Doppler shift the intrinsic signal frequency of a binary CW signal can undergo due to orbital motion. From the phase model of Eq. (29) we see that the instantaneous Doppler shift is

$$\begin{aligned} \left| \frac{d\phi/dt}{2\pi f} - 1 \right| &= \left| \frac{dR}{c dt} \right| \\ &= a_p \Omega \left| \frac{\sqrt{1-e^2} \cos E \cos \omega - \sin E \sin \omega}{1-e \cos E} \right| \\ &\leq a_p \Omega \frac{|\sin E \sin \omega| + |\cos E \cos \omega|}{|1-e \cos E|}, \end{aligned} \quad (\text{B1})$$

where we used the fact that $|a+b| \leq |a|+|b|$ and $\sqrt{1-e^2} \leq 1$. We further observe that

$$\begin{aligned} |\cos E \cos \omega| + |\sin E \sin \omega| &= \\ \max\{|\cos(E+\omega)|, |\cos(E-\omega)|\} &\leq 1, \end{aligned} \quad (\text{B2})$$

and $|1-e \cos E| \geq 1-e$ to obtain

$$\left| \frac{d\phi/dt}{2\pi f} - 1 \right| \leq \frac{a_p \Omega}{1-e}. \quad (\text{B3})$$

Appendix C: Maximal SFT length

By using the SFT-based demodulation method to compute the $\tilde{\mathcal{F}}$ -statistic, the computing cost per template increases linearly with the number of SFTs used [see Eq. (88)]. On the other hand, the maximal length of the SFT is limited by the linear-phase approximation over the duration of each SFT. In order to minimize the computing cost of this method, we therefore want to choose the longest possible SFT duration T_{sft} [see Eq. (89)] with an acceptable error in the linear-phase approximation.

In order to estimate the maximal phase-error of the linear-phase approximation over an SFTs, we can conveniently re-use the short-segment limit expressions, and simply estimate the phase error as

$$|\Delta\phi| \sim |v_2| = 2\pi \frac{|u_2|}{2!} \left(\frac{T_{\text{sft}}}{2} \right)^2 \sim \frac{\pi}{4} a_p \bar{f} \Omega^2 T_{\text{sft}}^2, \quad (\text{C1})$$

and the corresponding mismatch is given via Eq. (57) as $\mu_{\text{sft}} \sim \tilde{g}_{22}^{\text{SS},v} v_2^2 = \frac{4}{45} v_2^2$. Turning this around to express the maximal T_{sft} for a given maximal mismatch μ_{sft} , we obtain

$$T_{\text{sft}}^2(\bar{f}) \leq \frac{6\sqrt{5}\mu_{\text{sft}}}{\pi a_p \bar{f} \Omega^2}, \quad (\text{C2})$$

where the *largest* values of the parameter space being searched should be used for a_p , Ω and \bar{f} . This constraint on T_{sft} is illustrated in Fig. 8 as a function of search frequency \bar{f} for Scorpius X-1 parameters of Table I. We

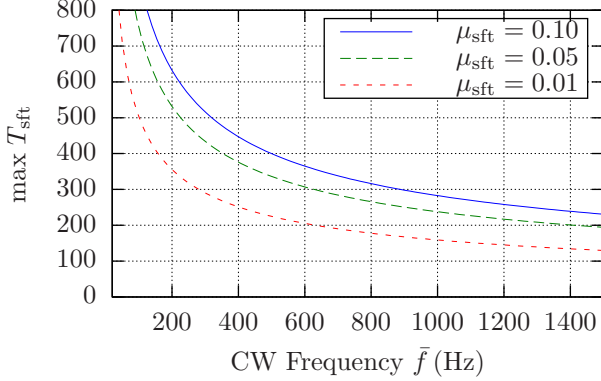


FIG. 8. Maximal allowed SFT length T_{sft} of Eq. (C2) for a Scorpius X-1 demodulation $\tilde{\mathcal{F}}$ -statistic search as function of frequency \bar{f} , for three different tolerated maximal mismatches μ_{sft} .

see that for Scorpius X-1 this limit is more stringent than the analogous constraint coming purely from the Doppler effect due to the detector motion, which typically results in a choice of $T_{\text{sft}} = 1800$ s in searches for isolated CW sources.

Appendix D: Optimal StackSlide solution for degenerate computing cost function

In the case of the FFT-based resampling algorithm for the $\tilde{\mathcal{F}}$ -statistic of Eq. (90), we encounter a degeneracy in the computing-cost function that had not been considered in the original optimization study of [20]. Namely, when the search dimension Ω is resolved in the semi-coherent template bank, but unresolved in the per-template coherent banks, then we see from Eqs. (72)–(75) that the number of templates (both for including eccentricity $\tilde{n} = 5, \hat{n} = 6$, and for the circular case $\tilde{n} = 3, \hat{n} = 4$), scale as $\tilde{N} \propto \Delta T$ and $\hat{N} \propto N \Delta T$, respectively, where we assumed the ideal gapless case with $\gamma = N$. From the computing-cost expressions Eq. (88) and Eq. (91) we see that therefore the computing-cost functions take the form $\tilde{\mathcal{C}} \propto N \Delta T \propto T_{\text{obs}}$ and $\hat{\mathcal{C}} \propto (N \Delta T)^2 \propto T_{\text{obs}}^2$. In this case the power-law coefficients in the formalism of [20] are $\tilde{\eta} = \tilde{\delta} = 1$ and $\hat{\eta} = \hat{\delta} = 2$ (the resulting N coefficients at fixed T_{obs} are therefore $\tilde{\varepsilon} = \hat{\varepsilon} = 0$), which corresponds to a degenerate case that has not been analyzed previously. In this case one can only find a solution by constraining T_{obs} . However, maximization of sensitivity at fixed computing cost would then result in $N \rightarrow 1$, i.e. a fully coherent search (except if by increasing ΔT the template bank eventually starts to resolve Ω and therefore breaks the degeneracy). In cases such as the Scorpius X-1 search considered here, there is an astrophysically-motivated upper bound on the segment length $\Delta T \leq 10$ d, and we therefore need to express the optimal solution for constraints on both T_{obs} and ΔT . With N and T_{obs} fixed, we can only optimize sensitivity over the respective template-bank mismatches $\tilde{\mu}_{\text{max}}$ and $\hat{\mu}_{\text{max}}$. This is achieved simply by solving Eq. (91) of [20], i.e. $(\tilde{\mu}_{\text{max}}/\tilde{n})/(\hat{\mu}_{\text{max}}/\hat{n}) = \tilde{\mathcal{C}}/\hat{\mathcal{C}}$, together with $\tilde{\mathcal{C}} + \hat{\mathcal{C}} = C_0$ for $\tilde{\mu}_{\text{max}}$ and $\hat{\mu}_{\text{max}}$.

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